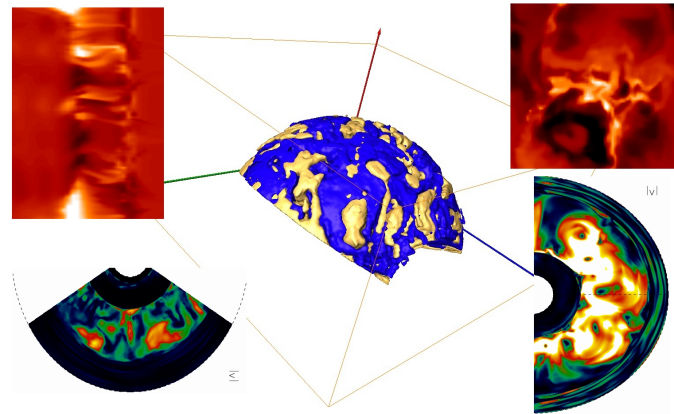


Experimentation with hydrogen injection flash

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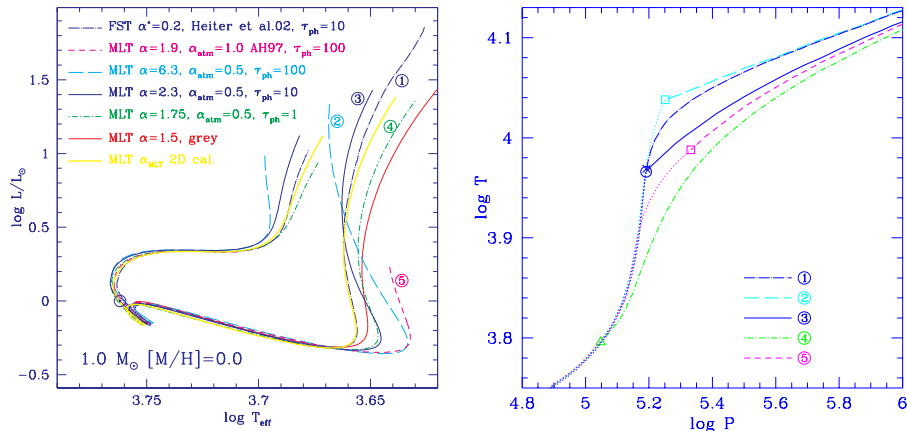


Outline

- Introduction
 - Why hydrodynamics?
- Simulations
 - Turbulent Entrainment
 - Entropy Barriers
 - Dual Core Flash
- Conclusions

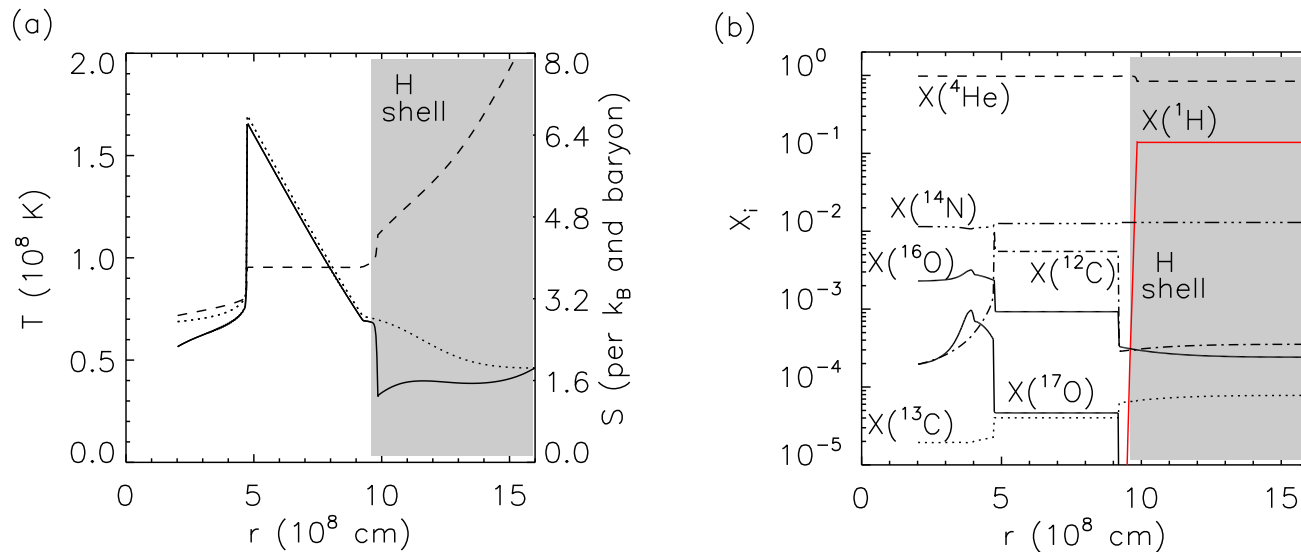
Why hydrodynamics?

- canonical stellar evolutionary calculations are hydrostatic and approximate dynamic processes by local and linear theories utilizing free parameters → **computationally not demanding** → produce observables
- hydrodynamic simulations are the best way to constrain the 1D approach as they are based on solution of Navier-Stokes equations and are essentially parameter free → **computationally demanding** → do not produce observables



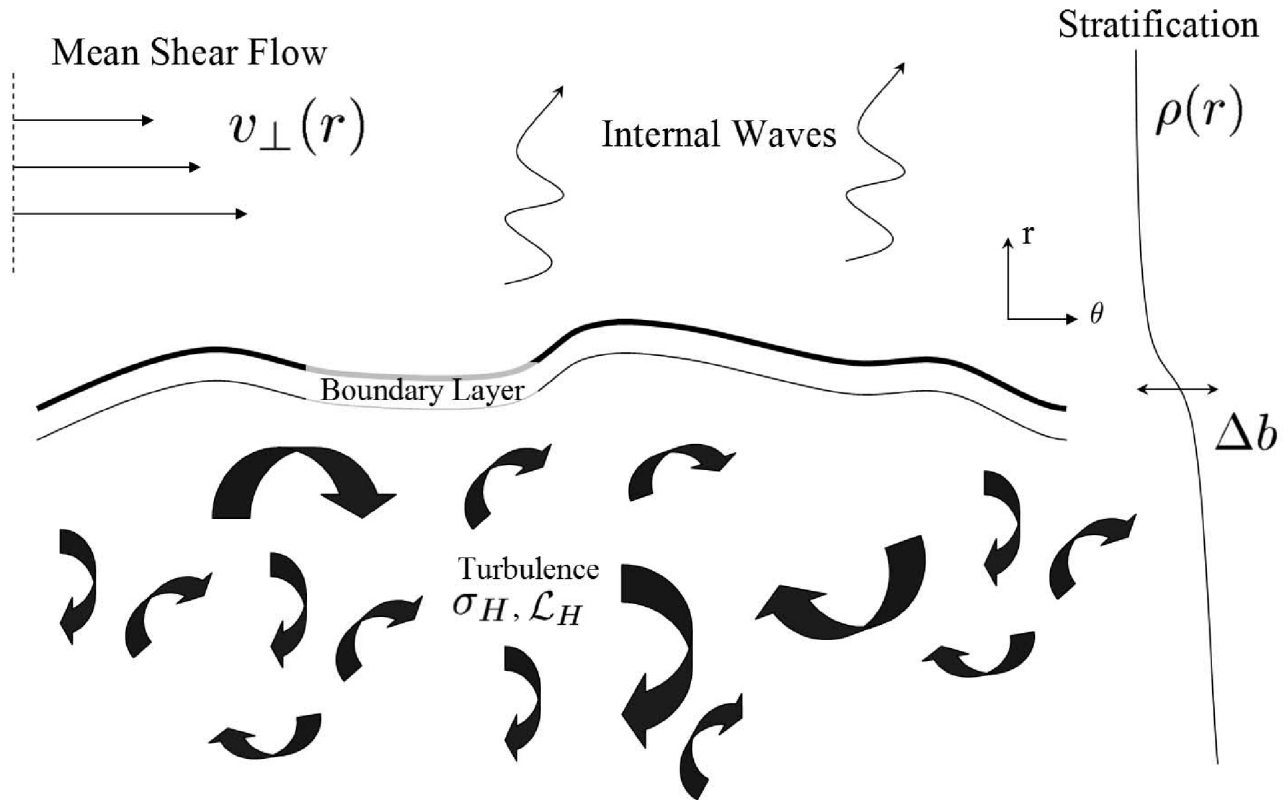
- Montalbán, J., D'Antona F., Kupka F., Heiter U., 2004, A&A, 416, 1081

Models

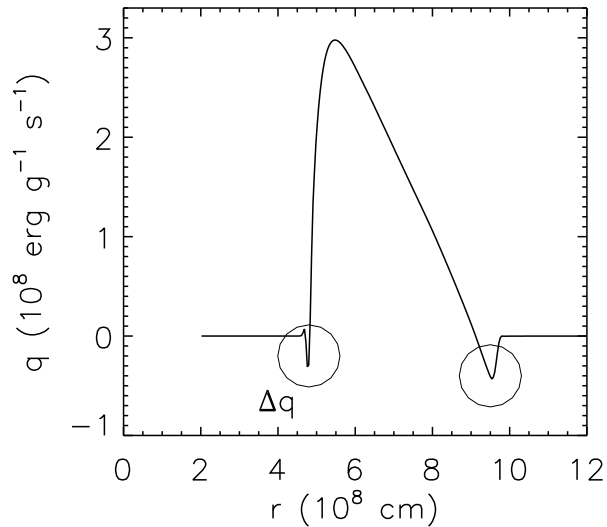
(a) Core Helium Flash $1.25 M_{\odot}$ with shifted hydrogen shell

- → watch the 3D movie

Turbulent Entrainment

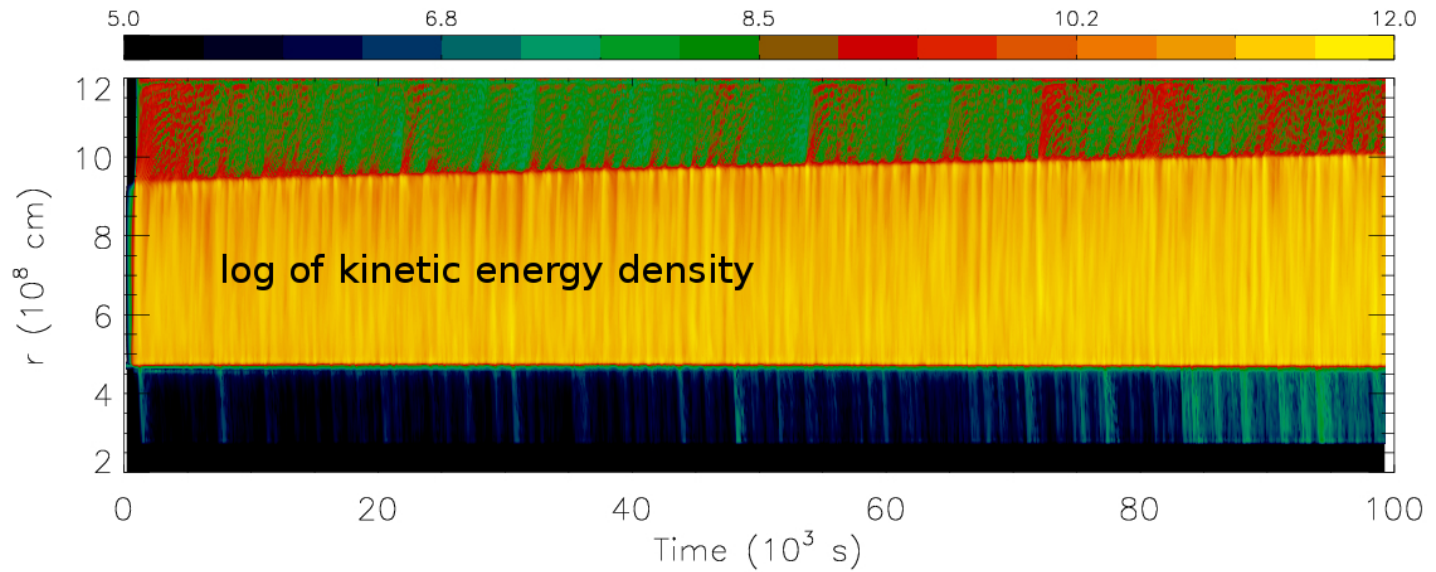


Turbulent Entrainment



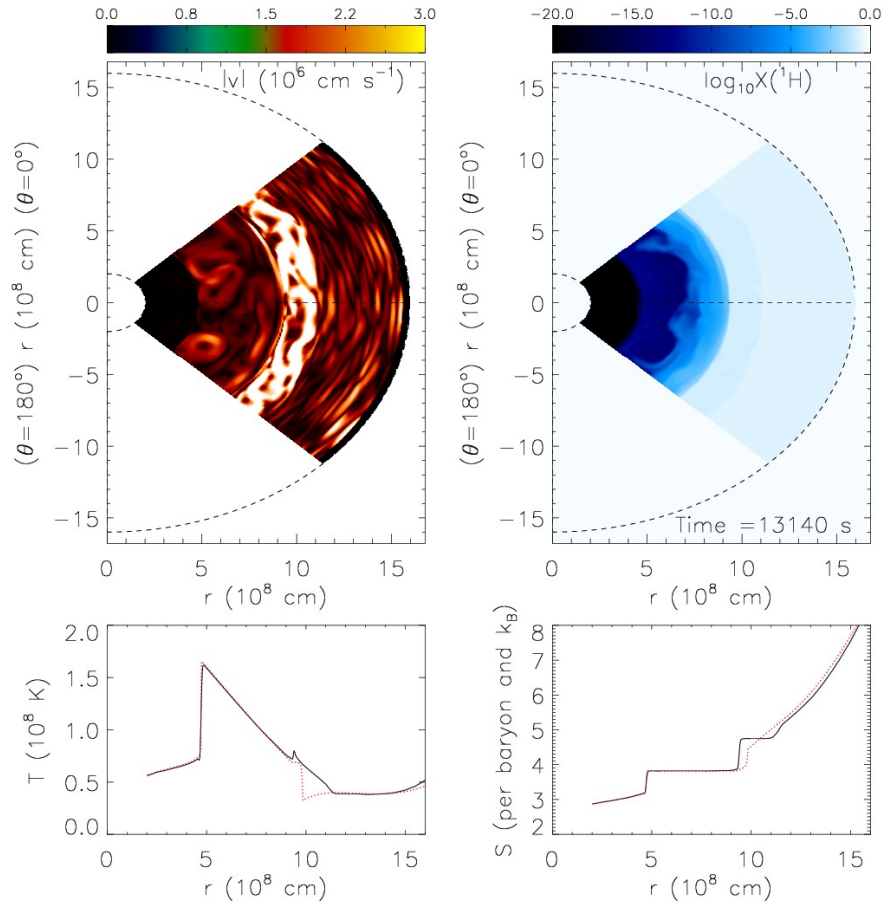
- accounts essentially for both, penetration and overshooting
- is an exchange between the potential energy of the stratification (buoyancy jump $db = N^2 dr$) and kinetic energy of the turbulence (buoyancy flux q)

Turbulent Entrainment and the Core Helium Flash



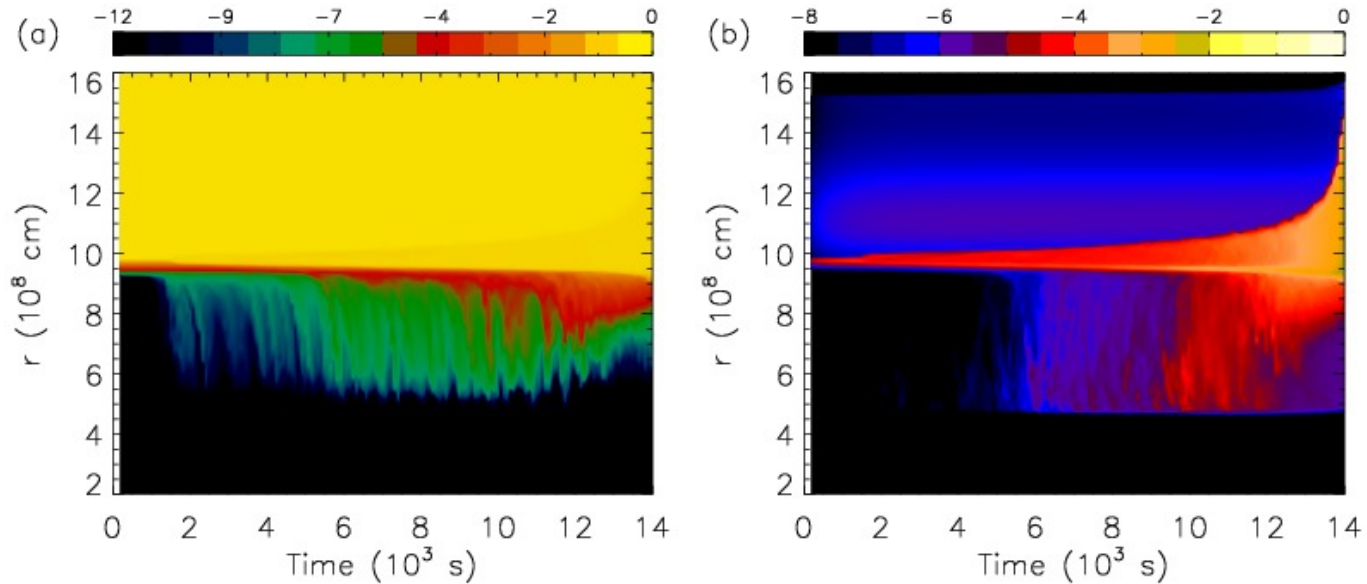
Mocák M., Campbell S. et al., Core Helium Flash Revisited III. From Pop I to Pop III stars (arXiv:1003.3646)

Is the entropy barrier at H-He discontinuity an ultimate obstacle?



- **NO, it is not** \rightarrow watch the movie

Mixing of elements during the hydrogen injection flash



(a) Core Helium Flash $1.25 M_{\odot}$ with shifted hydrogen shell

- mixing of elements during hydrogen injection flash from H shell to the He core not completely prohibited

Conclusions

- hydrogen injection flash is possible for low-mass stars of all metallicities
- entropy barriers not an ultimate obstacle for growing convection zones
- mixing of elements during hydrogen injection flash from H shell to the He core possible