



The **VORTEX** team

Université  
de Liège



*M. Reggiani, V. Christiaens, O. Absil, D. Mawet, E. Huby, E. Choquet, C. A. Gomez Gonzalez, G. Ruane, B. Femenia, E. Serabyn, K. Matthews, M. Barraza, B. Carlomagno, D. Defrere, C. Delacroix, S. Habraken, A. Jolivet, M. Karlsson, G. Orban de Xivry, P. Piron, J. Surdej, E. Vargas Catalan, O. Wertz*

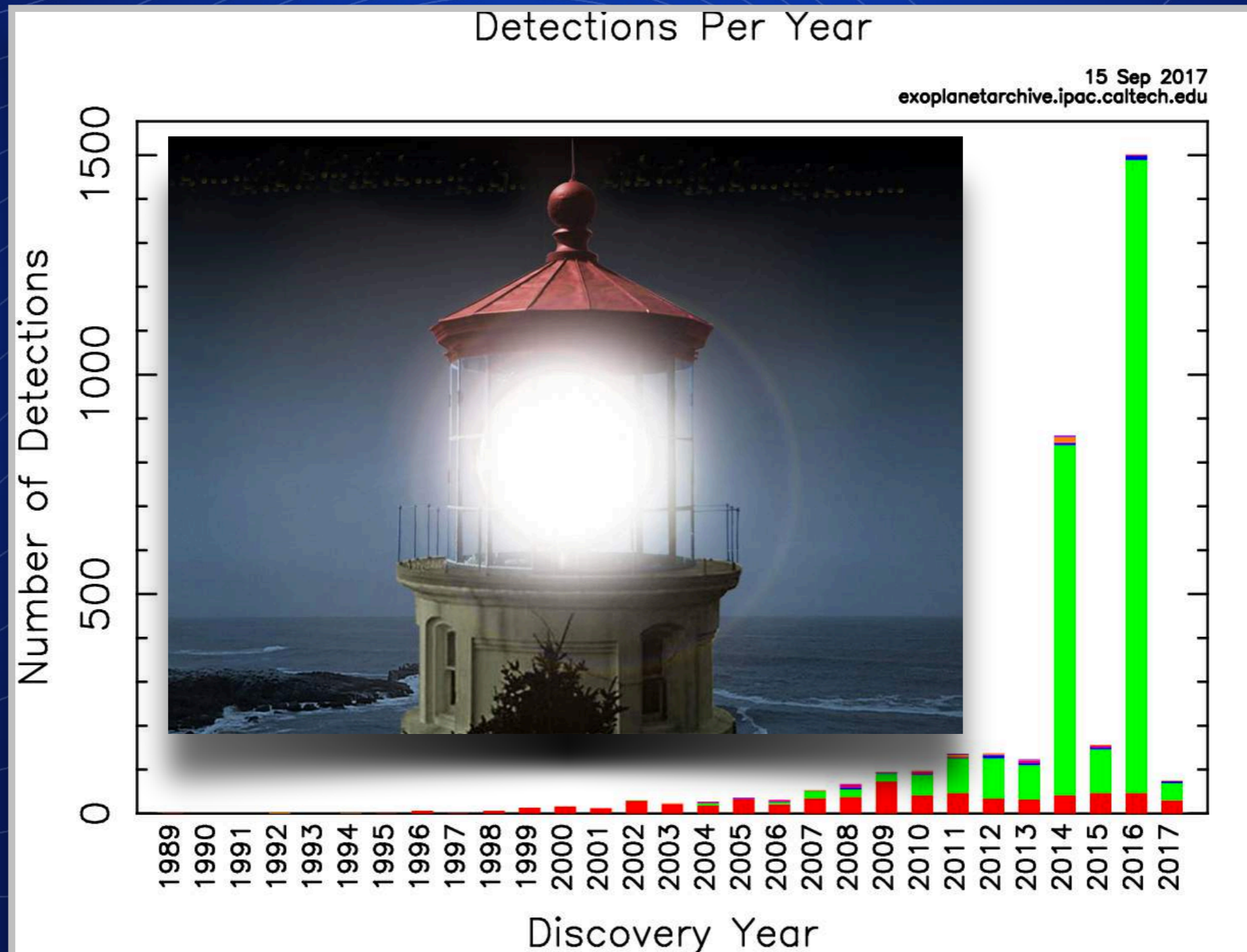
**MADDALENA REGGIANI & VORTEX TEAM**

---

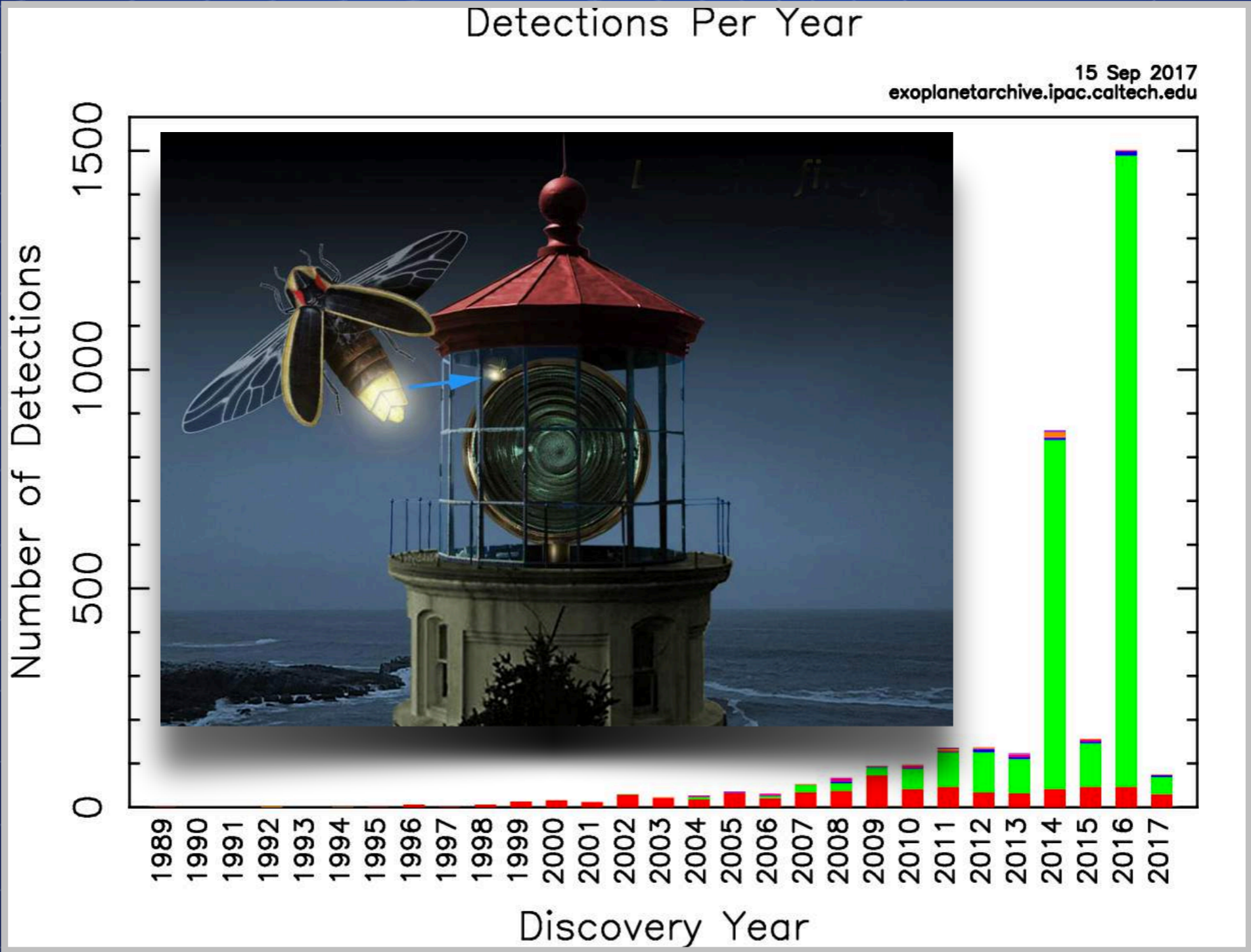
**A CANDIDATE ACCRETING PROTOPLANET IN  
THE TRANSITION DISK AROUND MWC 758**



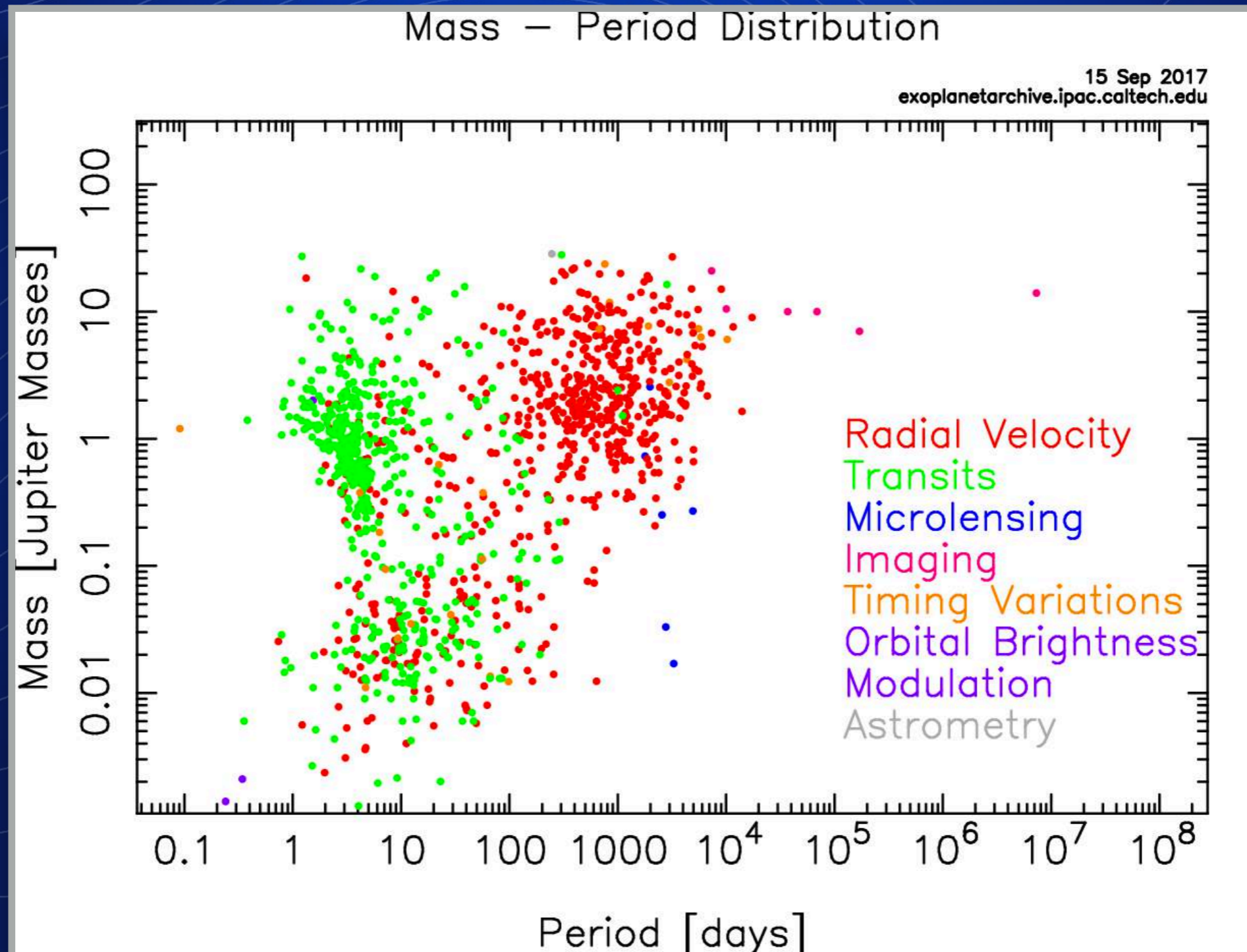
# DIRECT IMAGING OF EXOPLANETS: WHY?



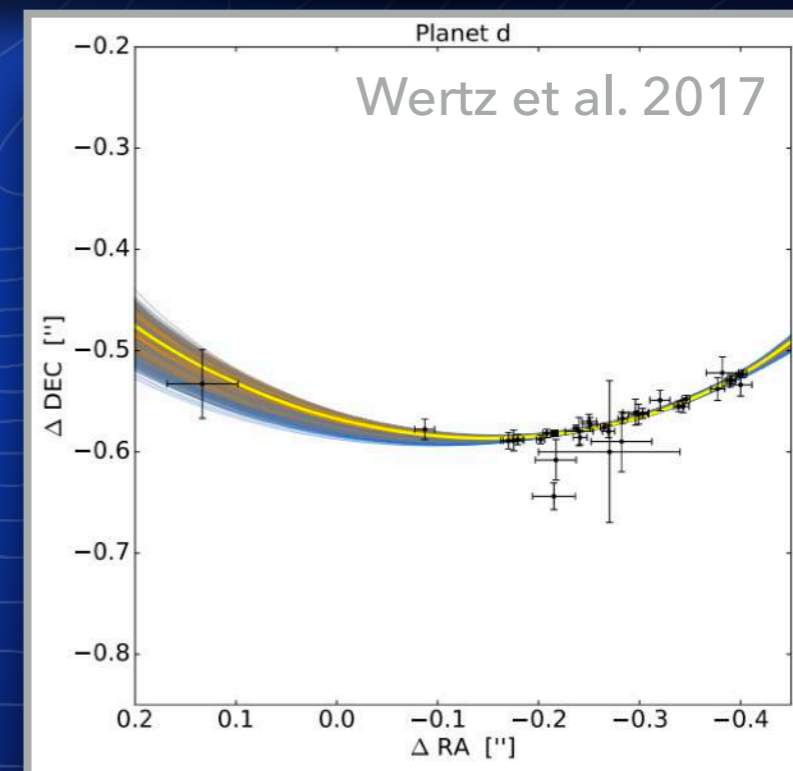
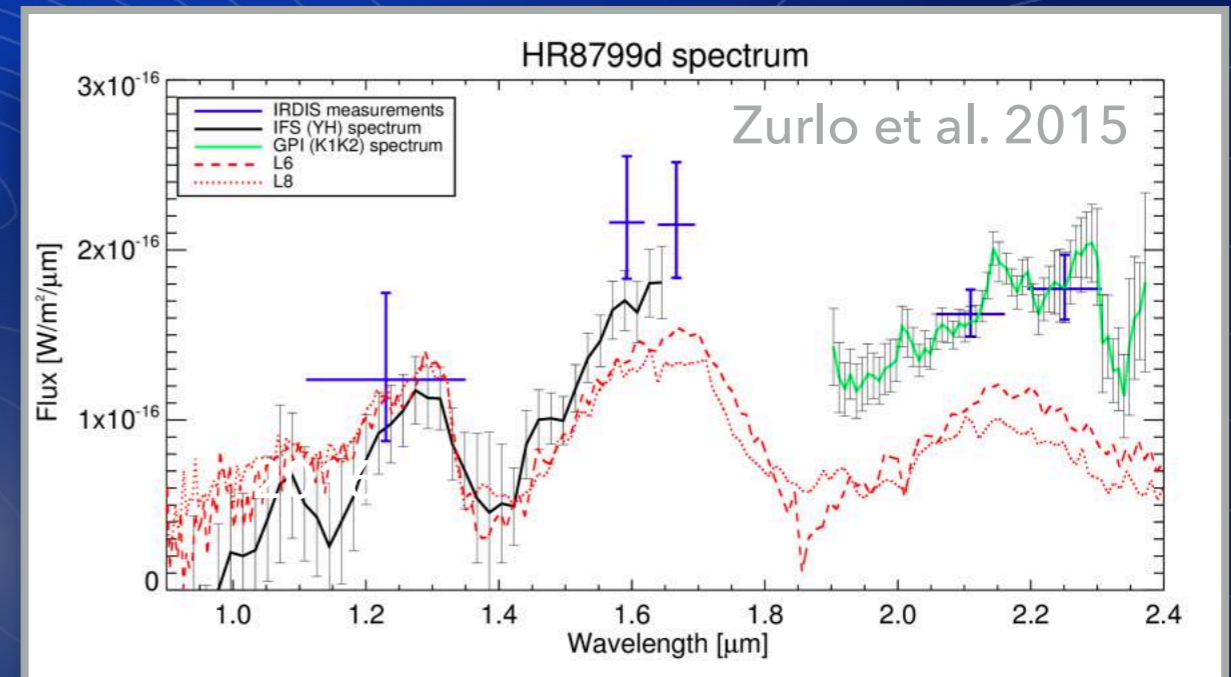
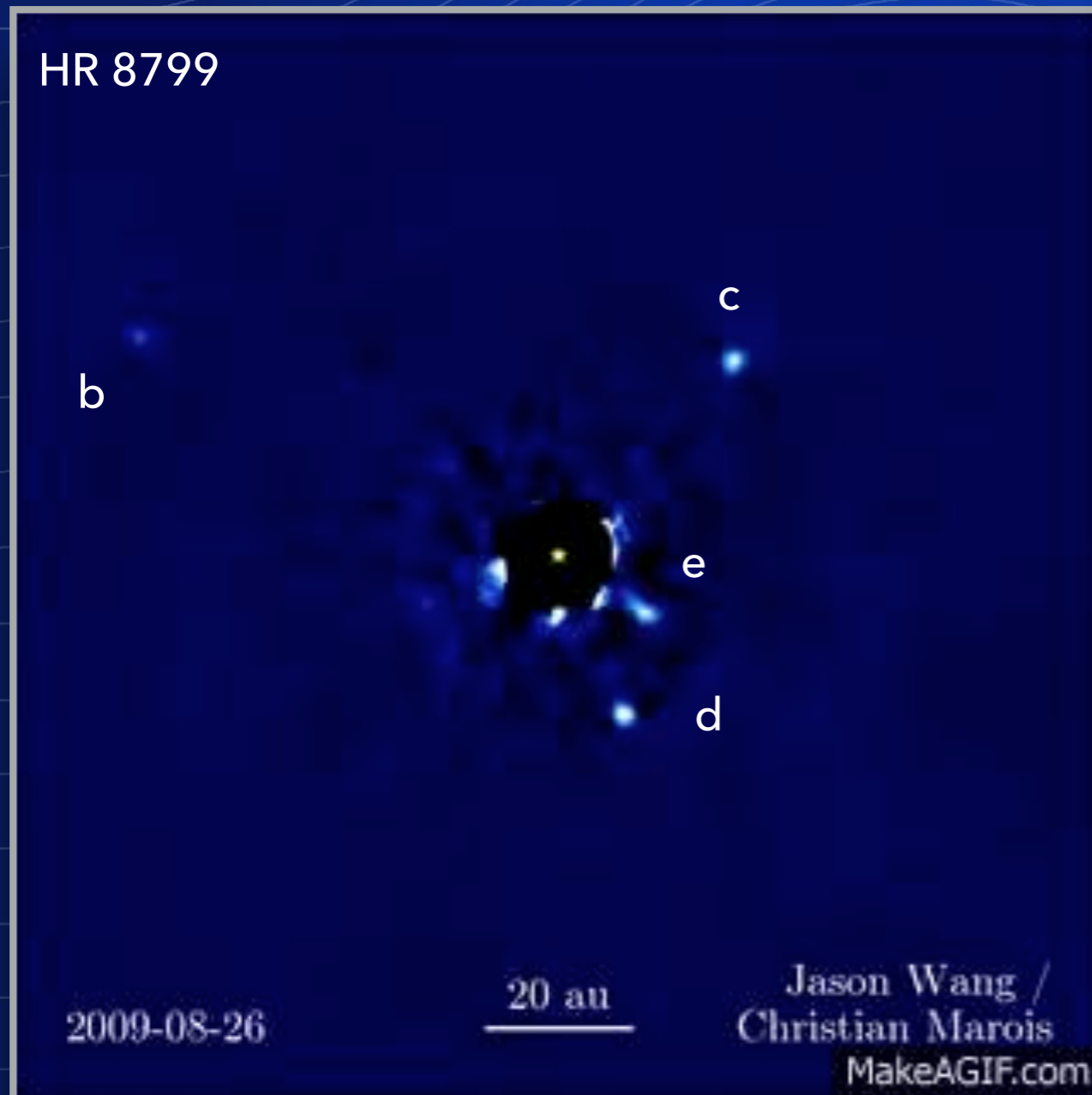
# DIRECT IMAGING OF EXOPLANETS: WHY?



# DIRECT IMAGING OF EXOPLANETS: WHY?



# DIRECT IMAGING OF EXOPLANETS: WHY?



# PLANET FORMATION IN A NUTSHELL

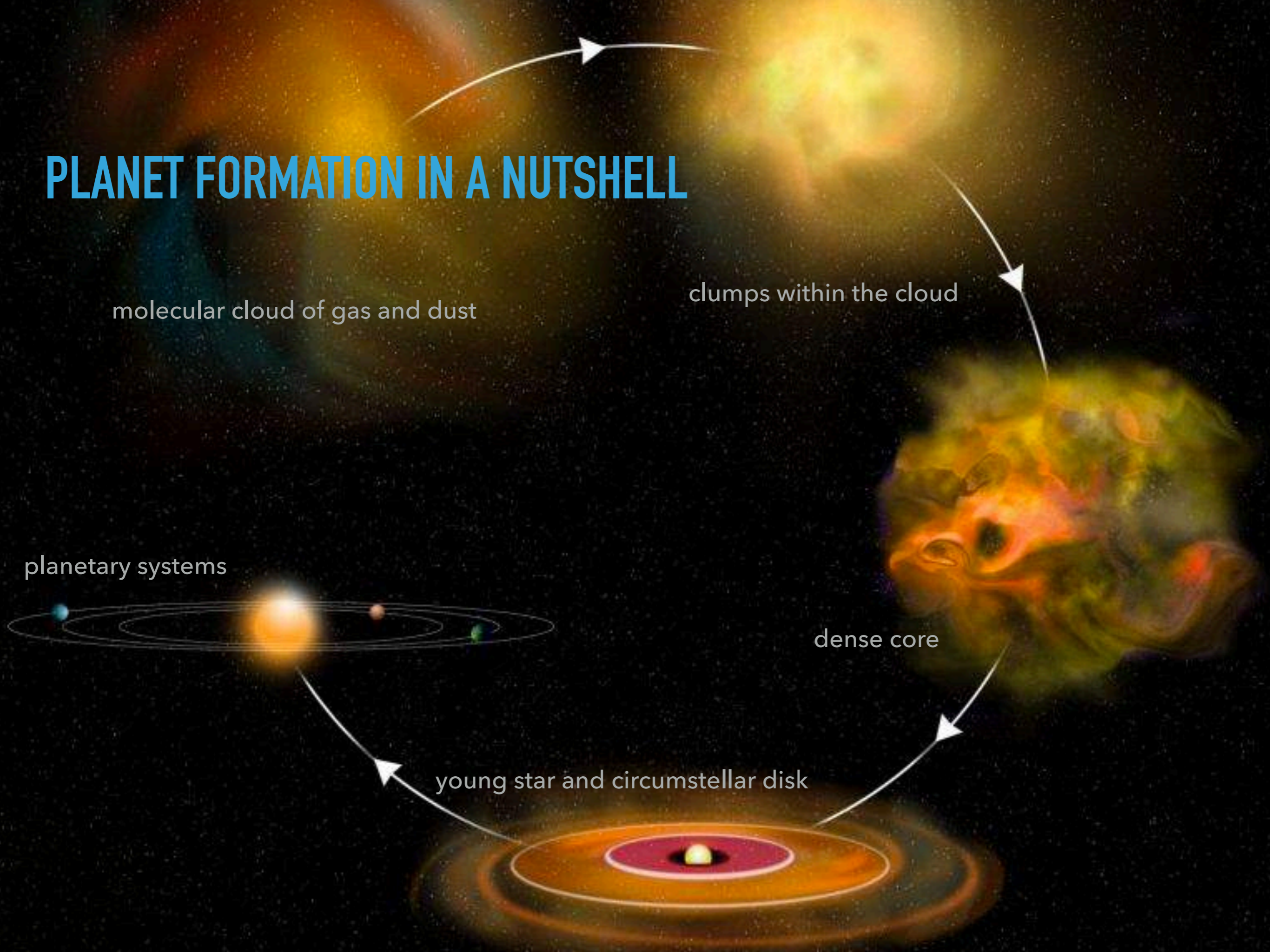
molecular cloud of gas and dust

clumps within the cloud

planetary systems

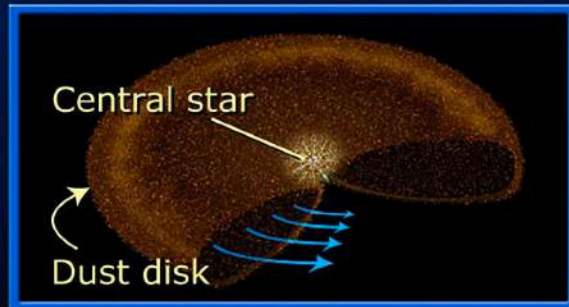
dense core

young star and circumstellar disk

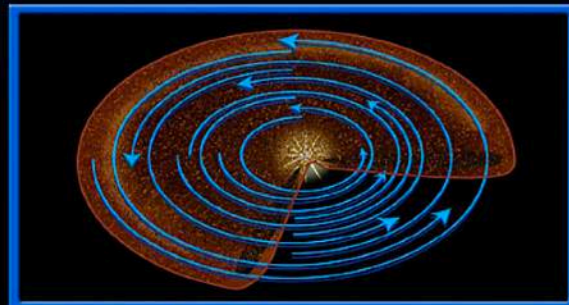


# TWO PLANET FORMATION SCENARIOS

## Accretion model



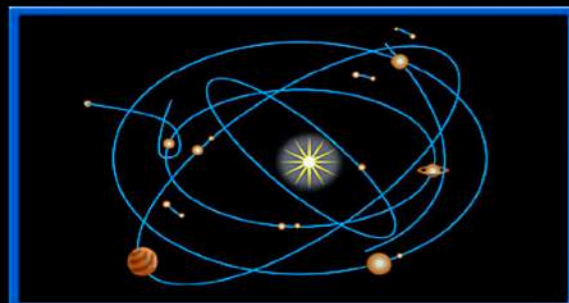
Orbiting dust grains accrete into "planetesimals" through nongravitational forces.



Planetesimals grow, moving in near-coplanar orbits, to form "planetary embryos."

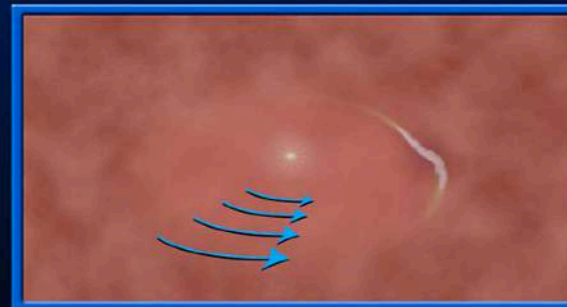


Gas-giant planets accrete gas envelopes before disk gas disappears.



Gas-giant planets scatter or accrete remaining planetesimals and embryos.

## Gas-collapse model



A protoplanetary disk of gas and dust forms around a young star.



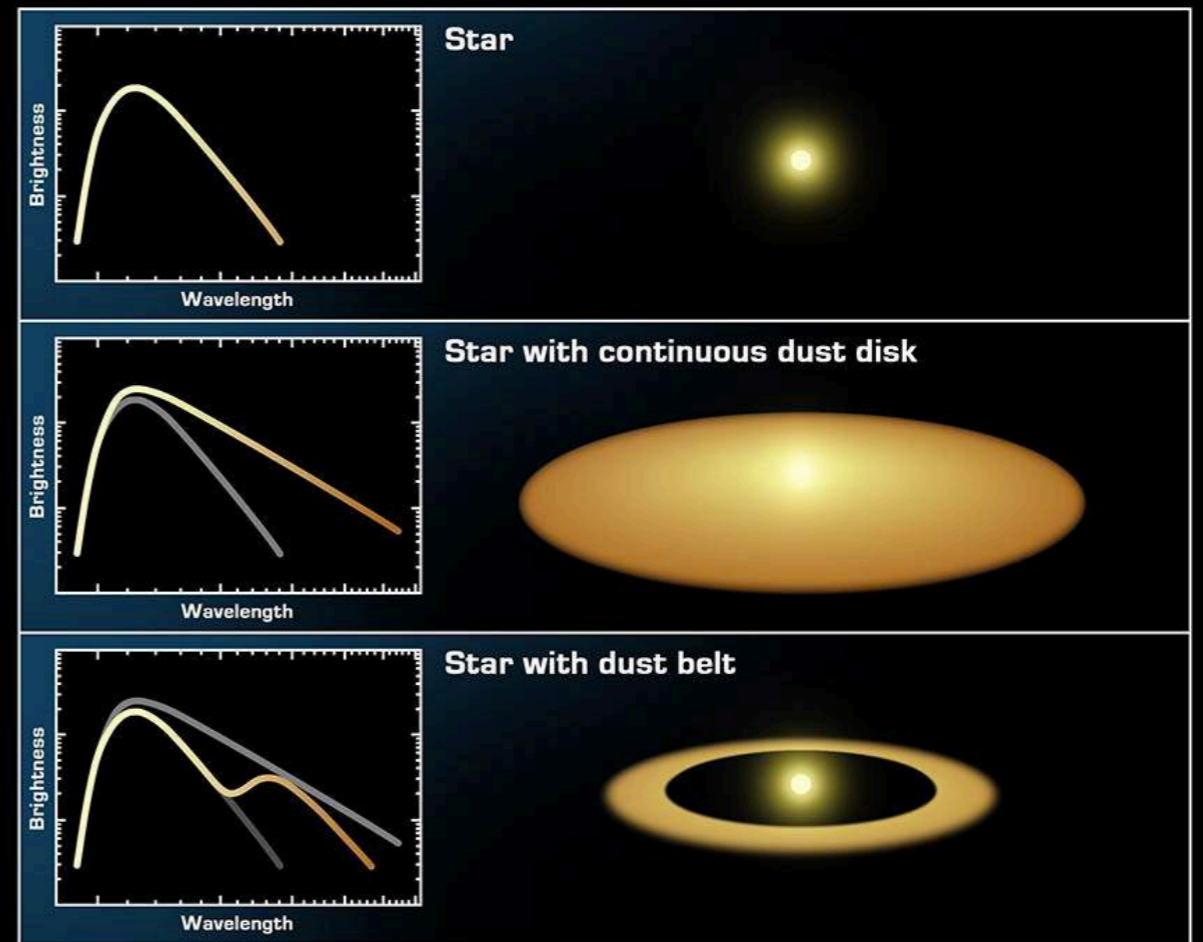
Gravitational disk instabilities form a clump of gas that becomes a self-gravitating planet.



Dust grains coagulate and sediment to the center of the protoplanet, forming a core.



The planet sweeps out a wide gap as it continues to feed on gas in the disk.

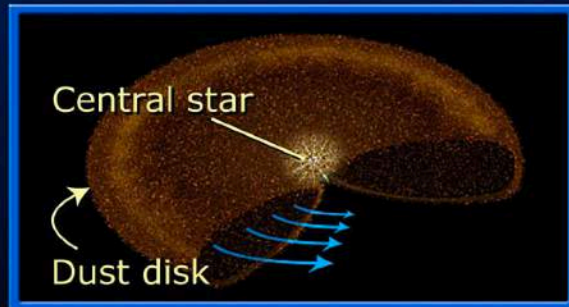


Giant planet formation must occur before the gas in the disk gets dissipated (max. ~10 Myr)

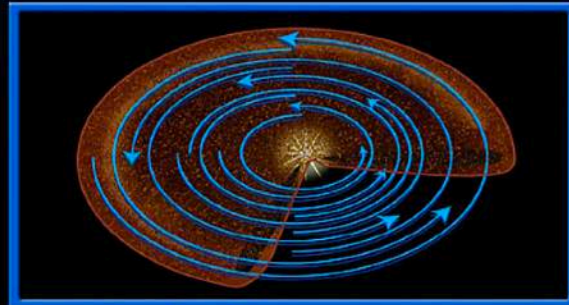


# TWO PLANET FORMATION SCENARIOS

## Accretion model



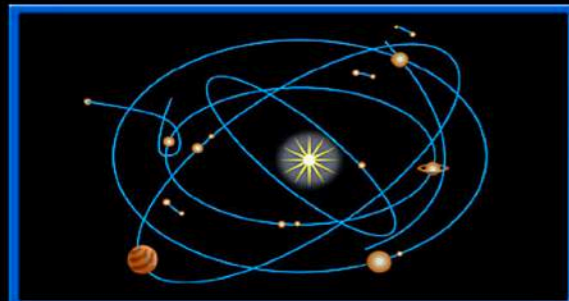
Orbiting dust grains accrete into "planetesimals" through nongravitational forces.



Planetesimals grow, moving in near-coplanar orbits, to form "planetary embryos."

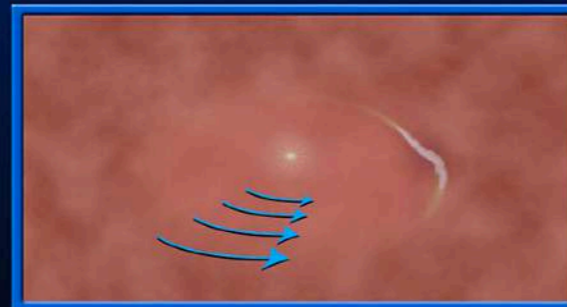


Gas-giant planets accrete gas envelopes before disk gas disappears.



Gas-giant planets scatter or accrete remaining planetesimals and embryos.

## Gas-collapse model



A protoplanetary disk of gas and dust forms around a young star.



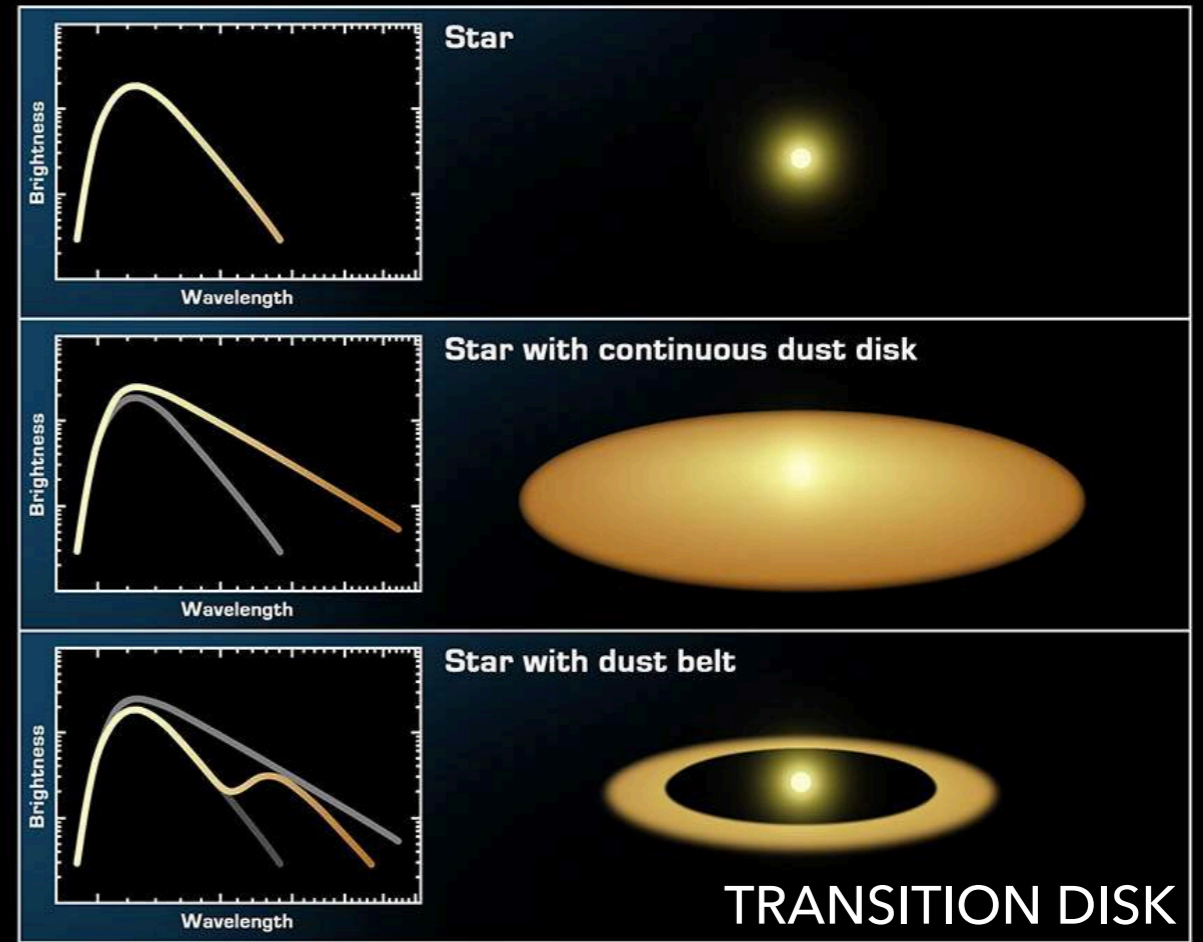
Gravitational disk instabilities form a clump of gas that becomes a self-gravitating planet.



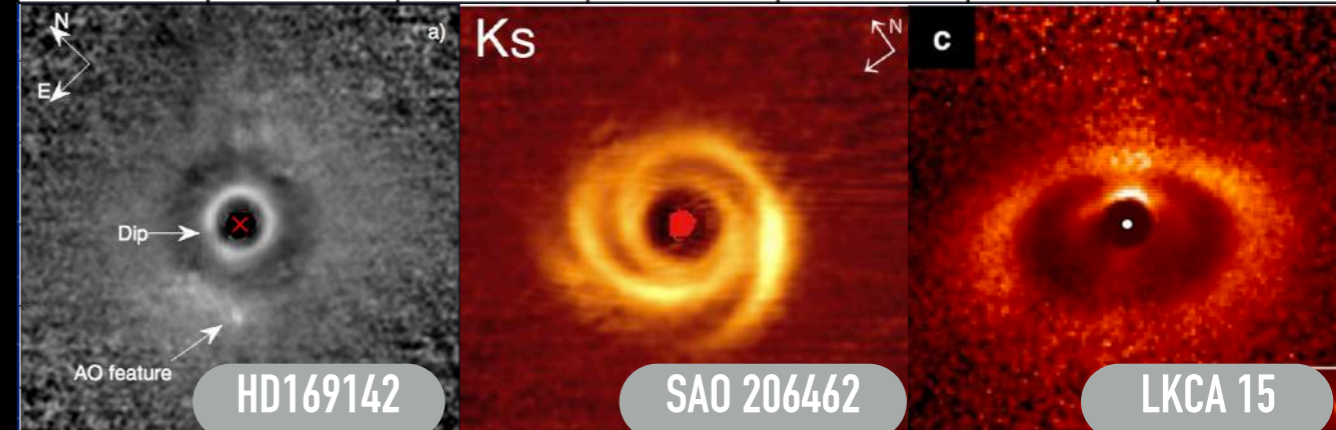
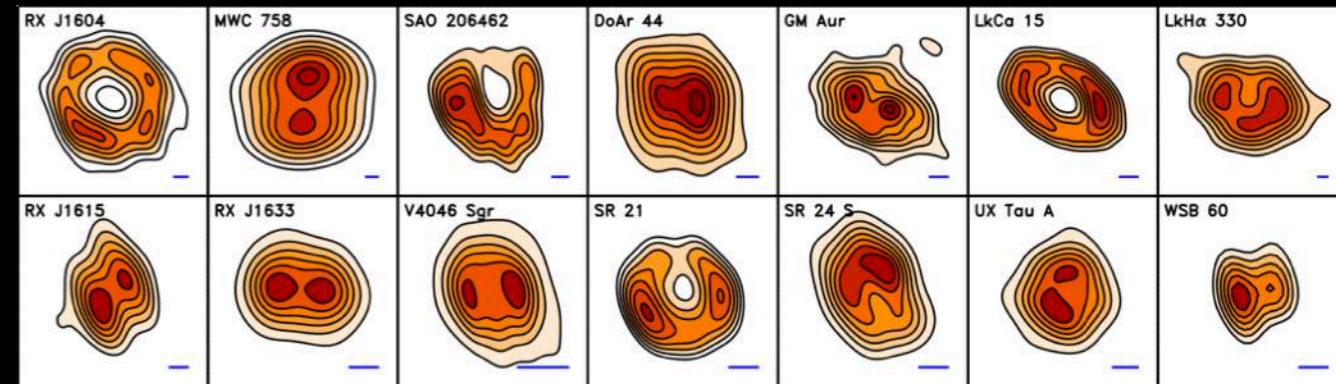
Dust grains coagulate and sediment to the center of the protoplanet, forming a core.



The planet sweeps out a wide gap as it continues to feed on gas in the disk.



Giant planet formation must occur before the gas in the disk gets dissipated (max. ~10 Myr)

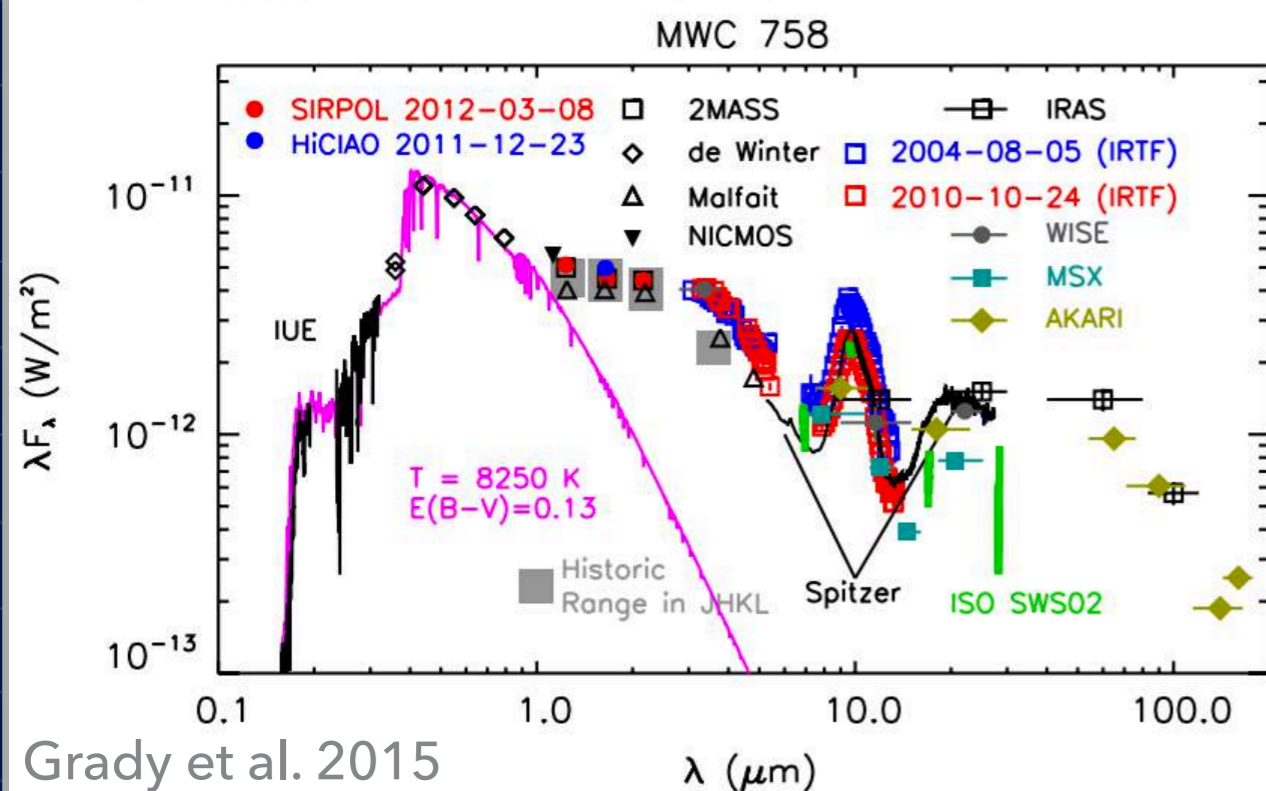
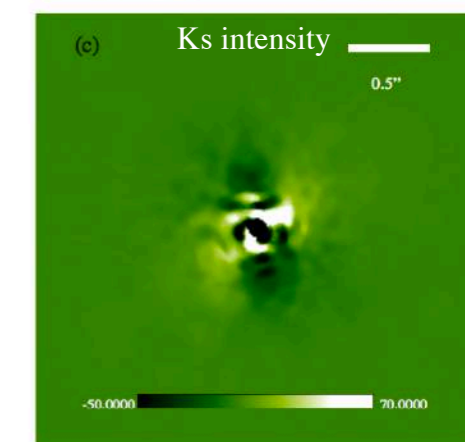
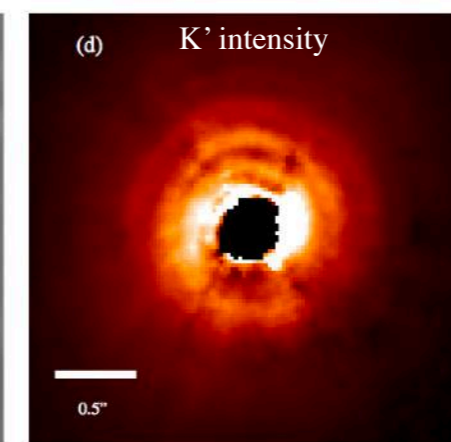
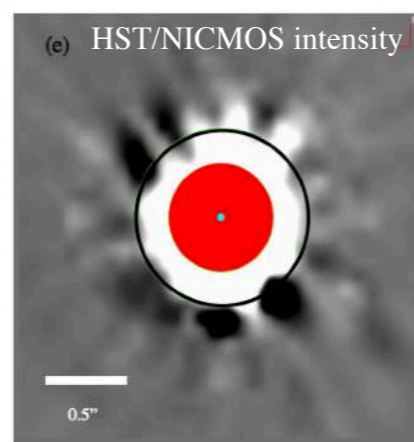
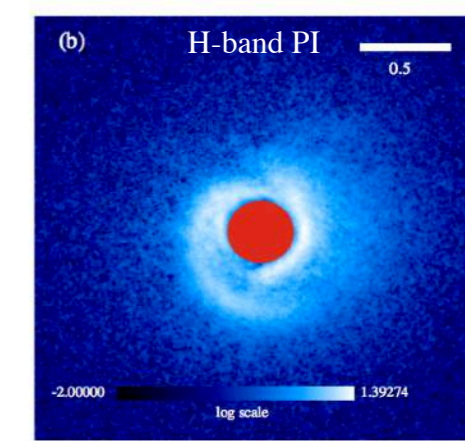
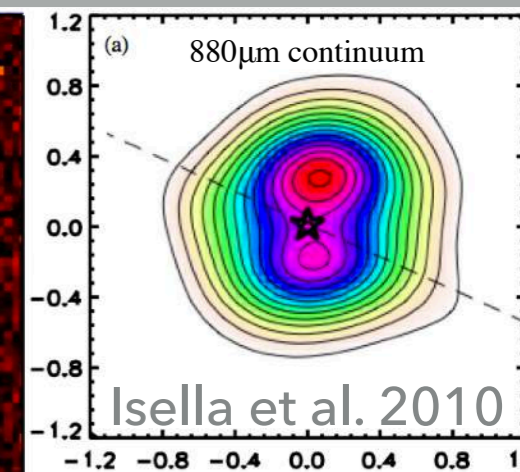
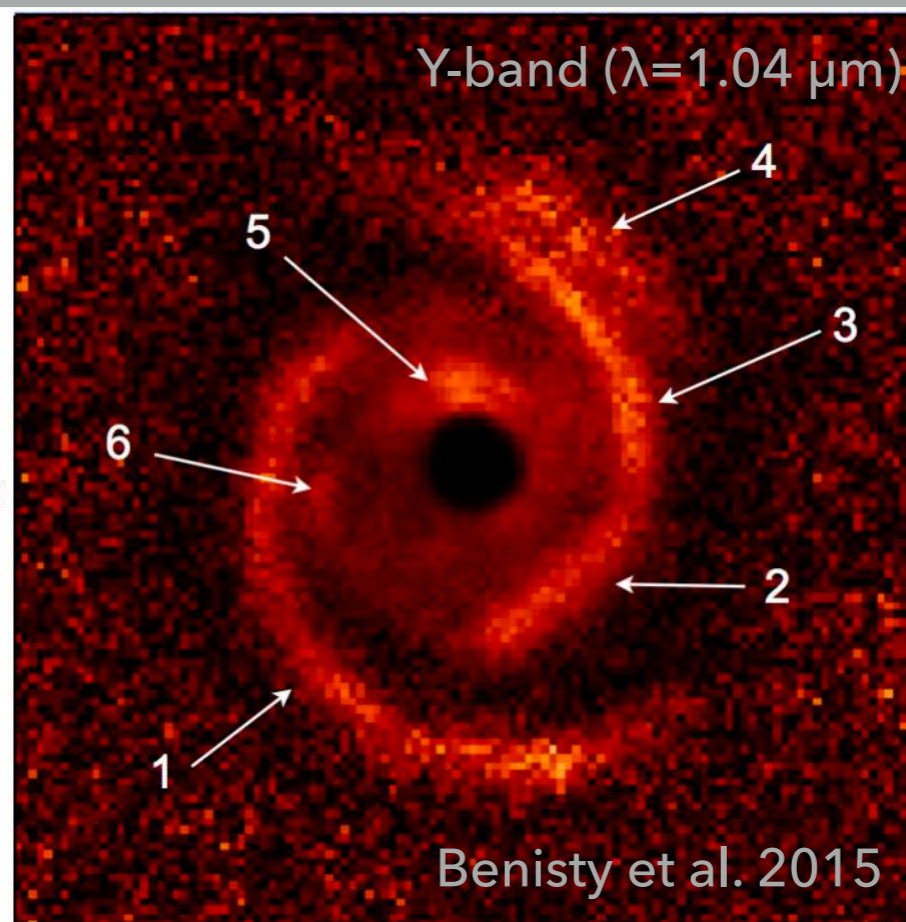


# TRANSITION DISK AROUND MWC 758

inclination: 21 degrees  
PA semi-major axis: 65 degrees

Properties	Values
RA (J2000)	05 <sup>h</sup> 30 <sup>m</sup> 27 <sup>s</sup> .530
DEC (J2000)	+25°19'57".082
Age (Myr)	3.5 ± 2.0 <sup>(1)</sup>
Mass (M <sub>⊙</sub> )	2.0 ± 0.2 <sup>(2)</sup>
L' (mag)	4.75 <sup>(3)</sup>
Distance (pc)	151 <sup>+8</sup> <sub>-9</sub> <sup>(4)</sup>

**References.** (1) Meeus et al. (2012); (2) Isella et al. (2010); (3) Malfait et al. (1998); (4) Gaia Collaboration (2016).

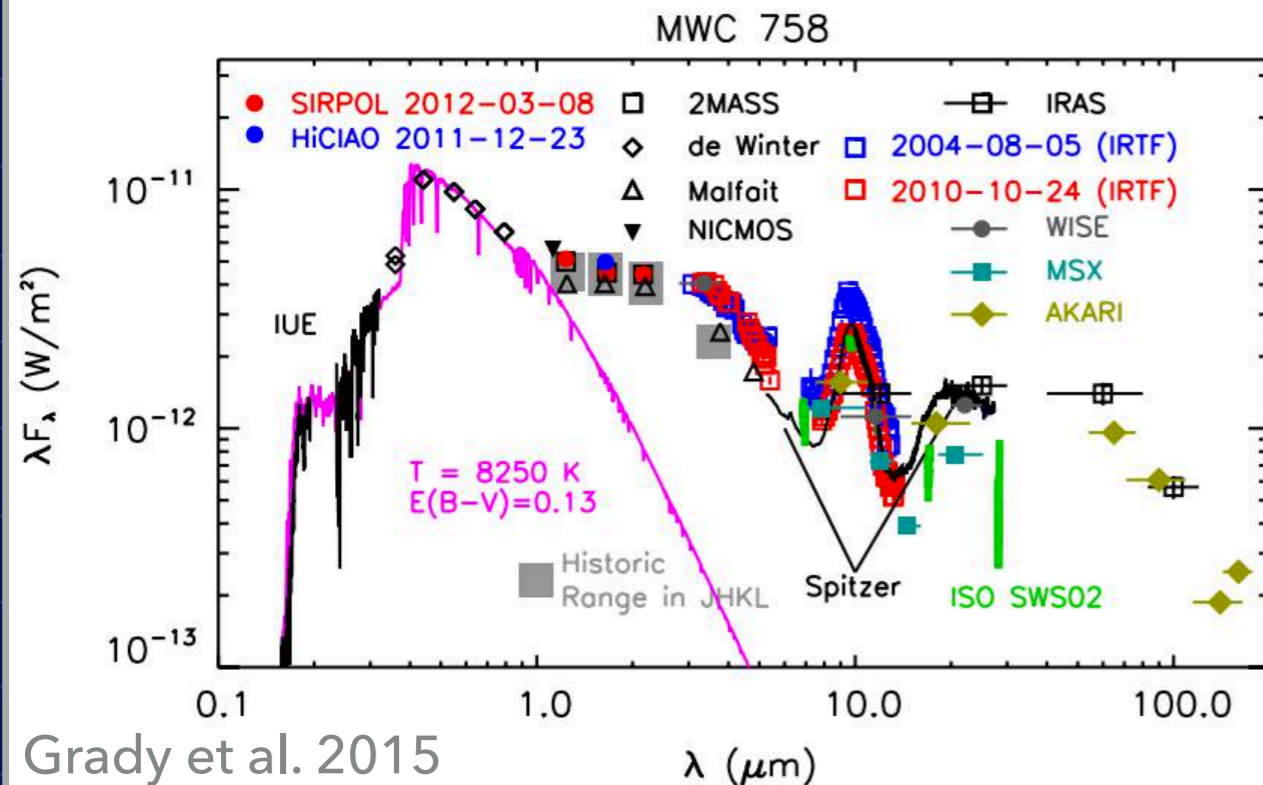
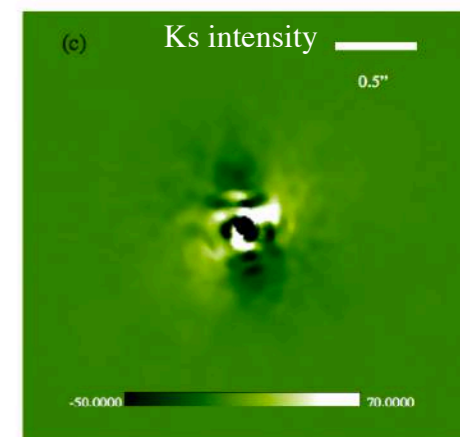
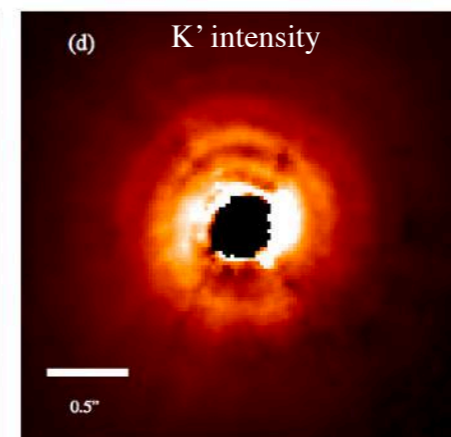
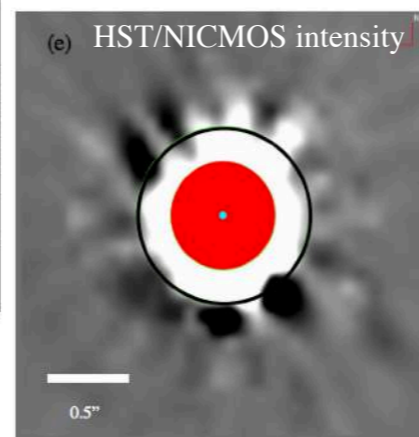
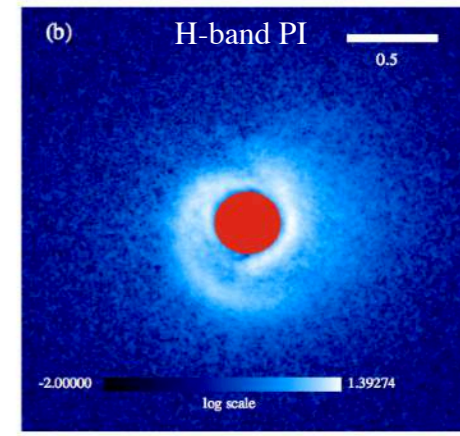
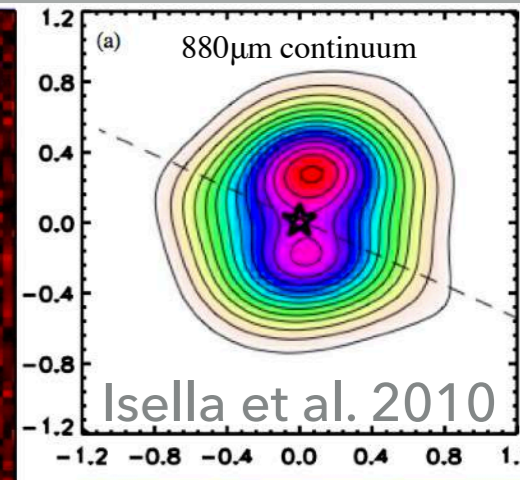
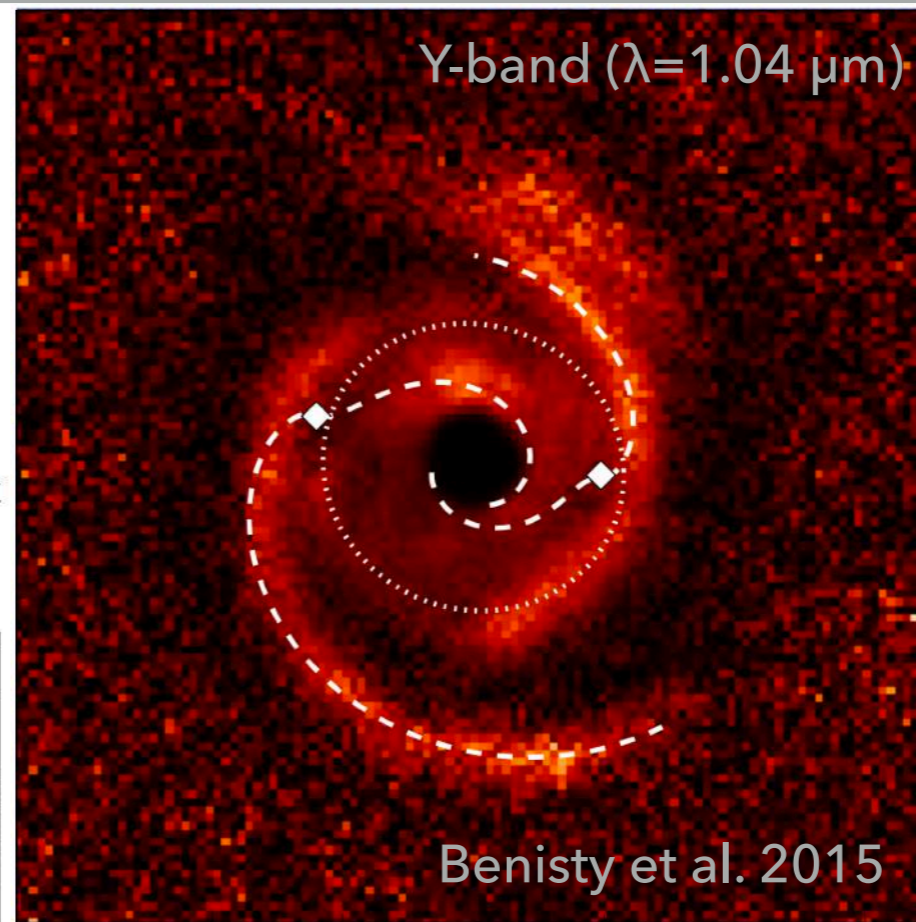


# TRANSITION DISK AROUND MWC 758

inclination: 21 degrees  
PA semi-major axis: 65 degrees

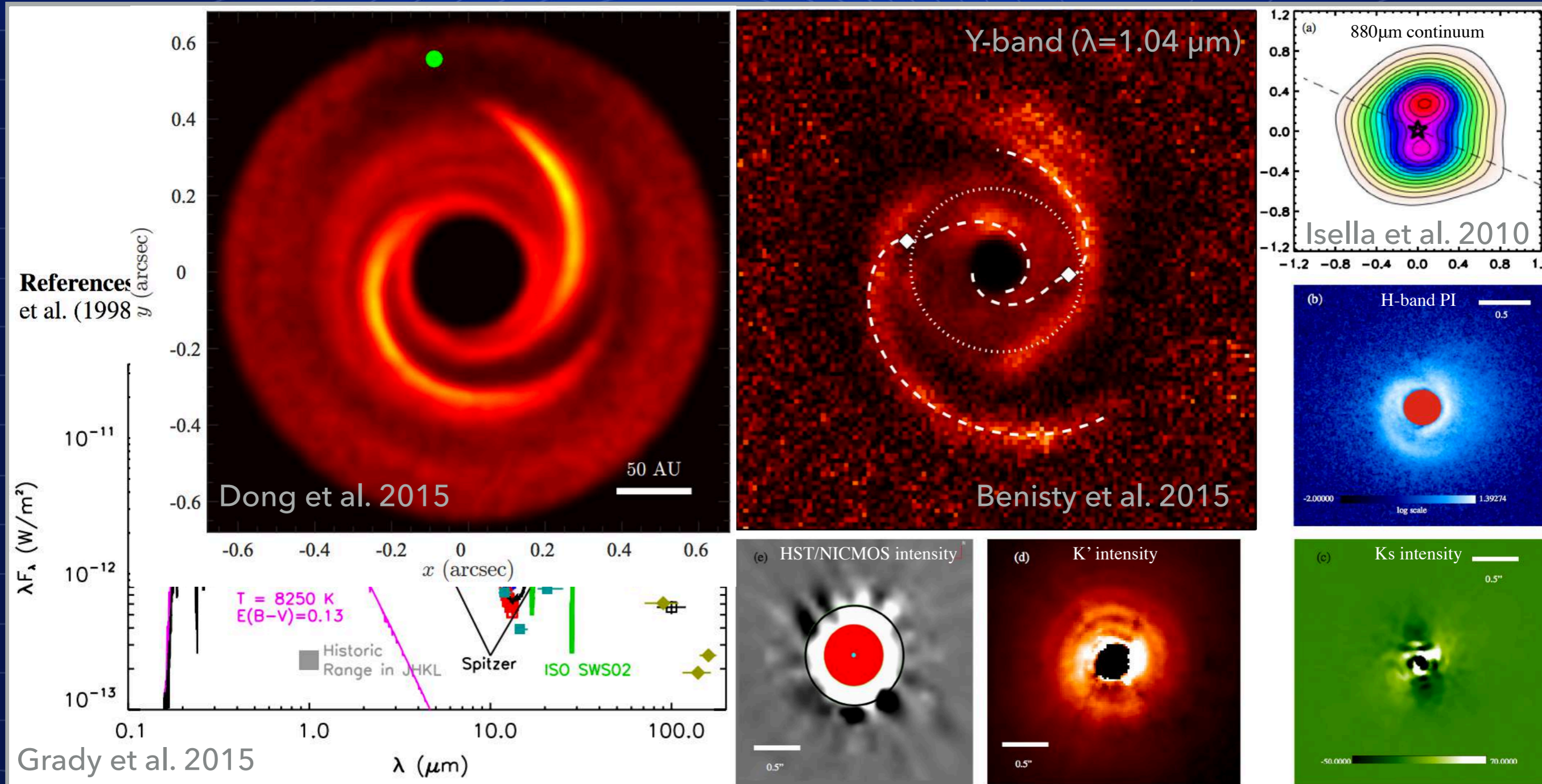
Properties	Values
RA (J2000)	05 <sup>h</sup> 30 <sup>m</sup> 27 <sup>s</sup> .530
DEC (J2000)	+25°19'57".082
Age (Myr)	3.5 ± 2.0 <sup>(1)</sup>
Mass (M <sub>⊙</sub> )	2.0 ± 0.2 <sup>(2)</sup>
L' (mag)	4.75 <sup>(3)</sup>
Distance (pc)	151 <sup>+8</sup> <sub>-9</sub> <sup>(4)</sup>

**References.** (1) Meeus et al. (2012); (2) Isella et al. (2010); (3) Malfait et al. (1998); (4) Gaia Collaboration (2016).



# TRANSITION DISK AROUND MWC 758

inclination: 21 degrees  
PA semi-major axis: 65 degrees



## KECK/NIRC2 OBSERVATIONS

- ▶ Keck/NIRC2 in ADI mode
- ▶ L' band ( $\lambda=3.8 \mu\text{m}$ )
- ▶ AGPM vortex coronagraph  
(Serabyn et al. 2017)
- ▶ operations fully automated with QUACITS: ensures consistent centering and data quality  
(Huby et al. 2015/2017)
- ▶ flat fielding, re-centering, bad pixel correction and bad frame removal, PCA-based background subtraction with VIP  
(Gomez Gonzalez et al. 2017)

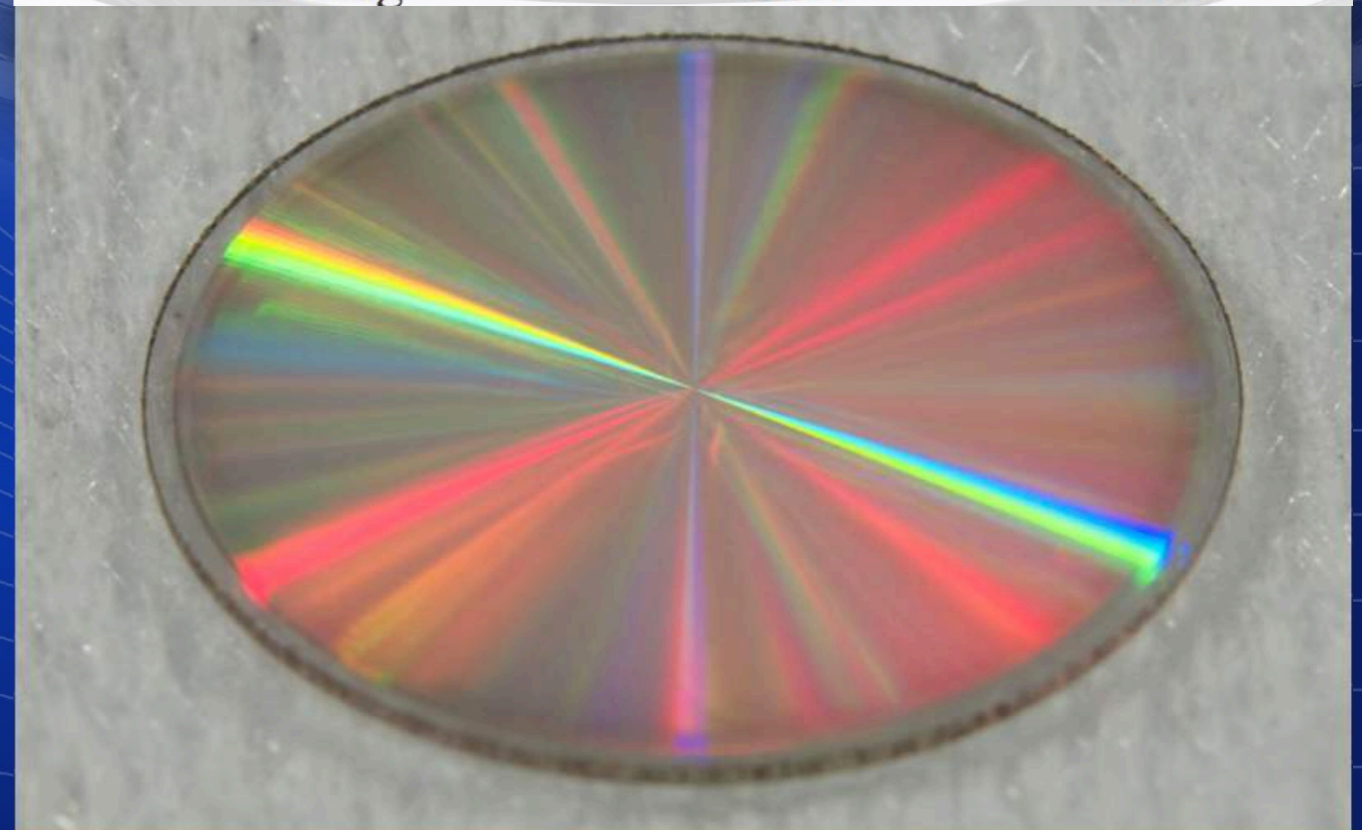
	First epoch	Second epoch
UT date (yyyy/mm/dd))	2015/10/24	2016/10/24
DIT (s)	0.5	0.25
Coadds	50	160
Number of frames	80	80
Total Int. Time (s)	2000	3200
Plate scale (mas/pix)	9.942	9.942
Filter Coronagraph	L'	L'
Par. angle start/end (°)	-128/+103	-90/+97
Mean airmass	1.012	1.074
Seeing	0.64	0.75



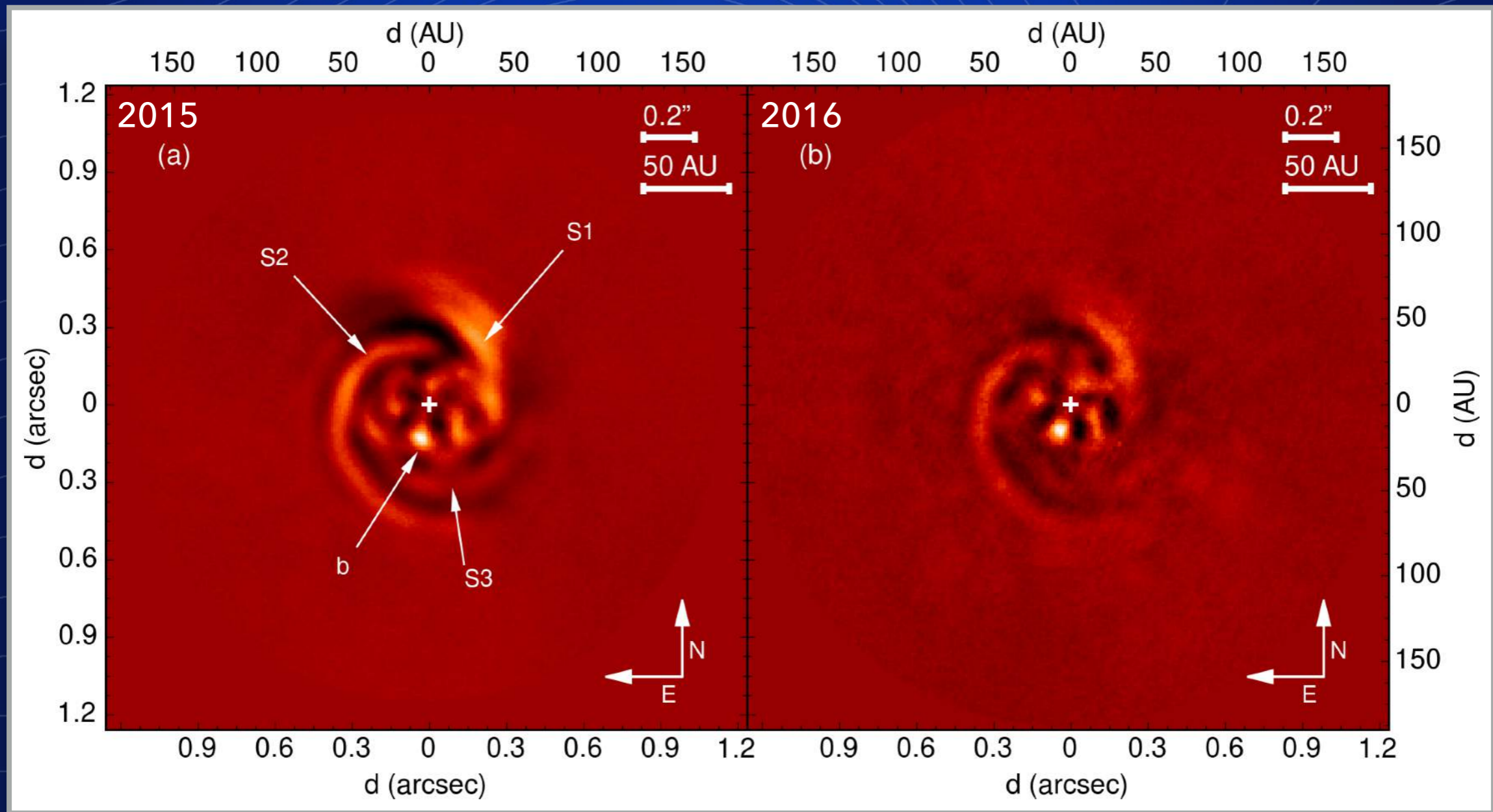
## KECK/NIRC2 OBSERVATIONS

- ▶ Keck/NIRC2 in ADI mode
- ▶ L' band ( $\lambda=3.8 \mu\text{m}$ )
- ▶ AGPM vortex coronagraph  
(Serabyn et al. 2017)
- ▶ operations fully automated with QUACITS: ensures consistent centering and data quality  
(Huby et al. 2015/2017)
- ▶ flat fielding, re-centering, bad pixel correction and bad frame removal, PCA-based background subtraction with VIP  
(Gomez Gonzalez et al. 2017)

	First epoch	Second epoch
UT date (yyyy/mm/dd))	2015/10/24	2016/10/24
DIT (s)	0.5	0.25
Coadds	50	160
Number of frames	80	80
Total Int. Time (s)	2000	3200
Plate scale (mas/pix)	9.942	9.942
Filter Coronagraph	L'	L'
Par. angle start/end (°)	-128/+103	-90/+97
Mean airmass	1.012	1.074
Seeing	0.64	0.75



# FINAL PCA/ADI IMAGES



Reggiani et al. (submitted)

# POINT-LIKE SOURCE: PROPERTIES AND EXPLANATIONS

UT date (yyyy/mm/dd)	$r^*$ (")	PA $^\ddagger$ ( $^\circ$ )	$\Delta L'^{\S}$ (mag)	$M_{L'}^{\#}$ (mag)
2015/10/24	$0.112 \pm 0.006$	$169 \pm 4$	$7.1 \pm 0.3$	11.85
2016/10/24	$0.110 \pm 0.006$	$162 \pm 5$	$6.9 \pm 0.5$	11.65

$R = 0.111 \pm 0.004$ ,  $\Delta L' = 7.0 \pm 0.3$  @  
DEPROJECTED SEPARATION OF  $20 \pm 1$  AU

▶ background source:

\* probability  $\approx 10^{-6}$  (TRILEGAL galactic model, Girardi et al. 2012 )

▶ asymmetric disk feature:

\* no maximum in polarized emission

▶ low-mass companion of 41-64 MJ (BT-SETTL models, Allard et al. 2012 ):

\* no fully depleted cavity in um-size dust  $\rightarrow$   $< 5.5$  MJ

( Benisty et al. 2015, Pinilla et al. 2015 )

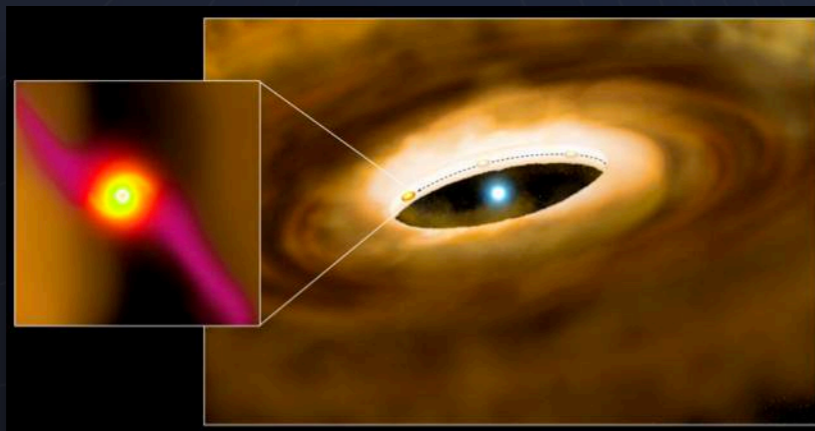


# POINT-LIKE SOURCE: AN ACCRETING PROTOPLANET, MWC 758 B

UT date (yyyy/mm/dd)	$r^*$ (")	PA <sup>‡</sup> (°)	$\Delta L'^{\S}$ (mag)	$M_{L'}^{\#}$ (mag)
2015/10/24	$0.112 \pm 0.006$	$169 \pm 4$	$7.1 \pm 0.3$	11.85
2016/10/24	$0.110 \pm 0.006$	$162 \pm 5$	$6.9 \pm 0.5$	11.65

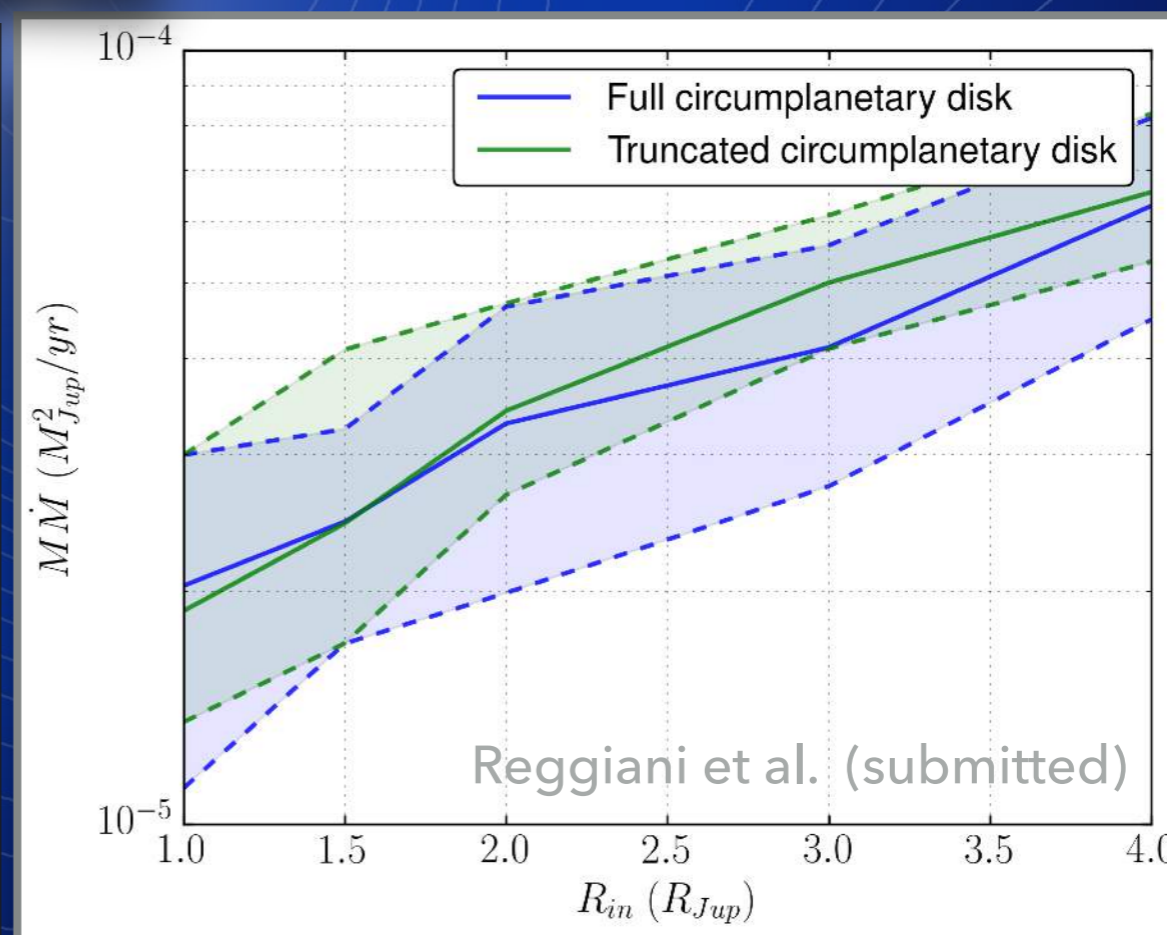
$R = 0.111 \pm 0.004$ ,  $\Delta L' = 7.0 \pm 0.3$  @  
DEPROJECTED SEPARATION OF  $20 \pm 1$  AU

▶ An accreting protoplanet,  
MWC 758b



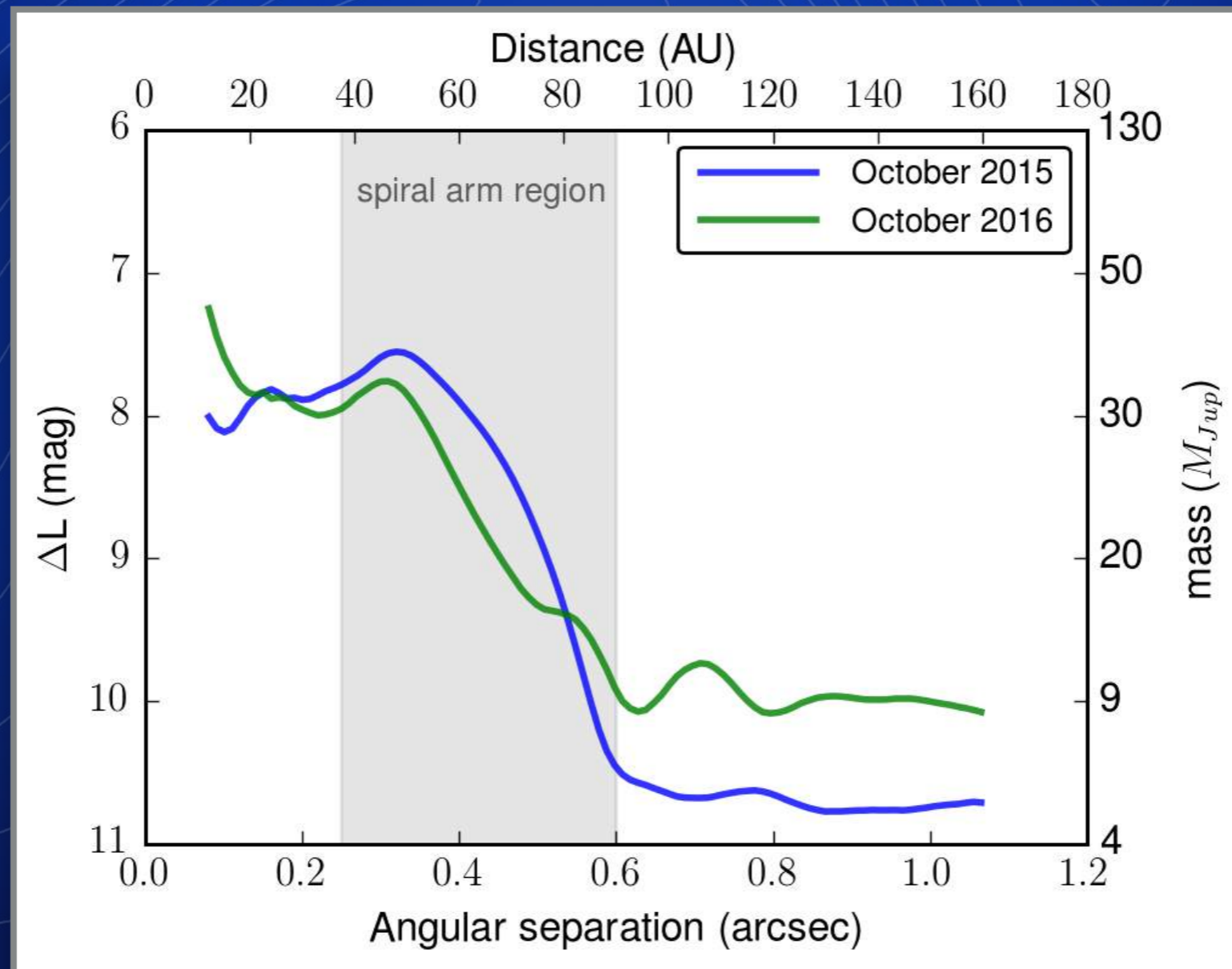
\*  $0.5-5 M_J$  @  $10^{-7}-10^{-9} M_{\odot}/yr$

(from Zhu et al. 2015 accretion models)



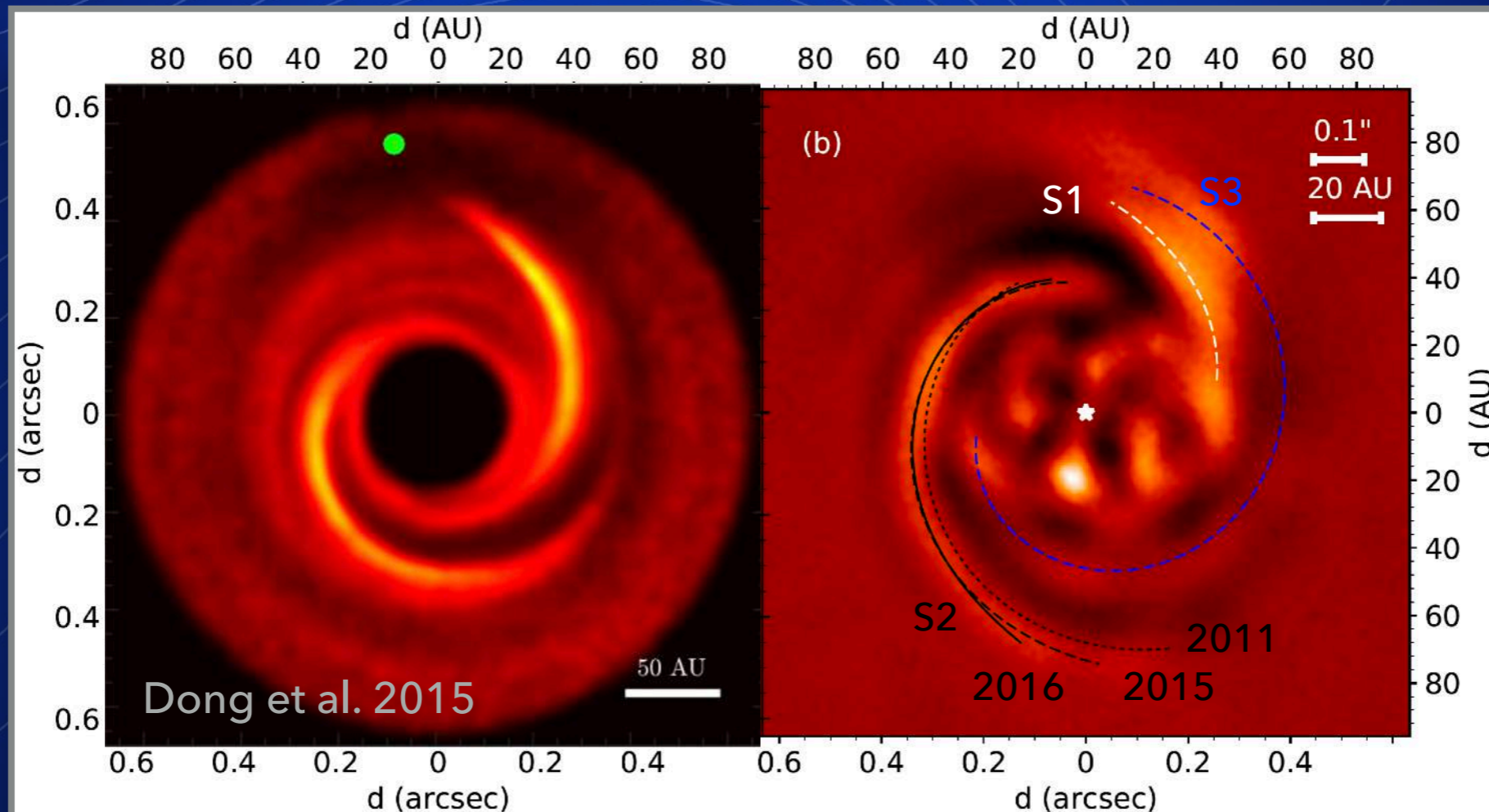
# CONSTRAINTS ON OTHER COMPANIONS

Detection limits at 95% completeness,  
following Ruane et al. 2017 :



Reggiani et al. (submitted)

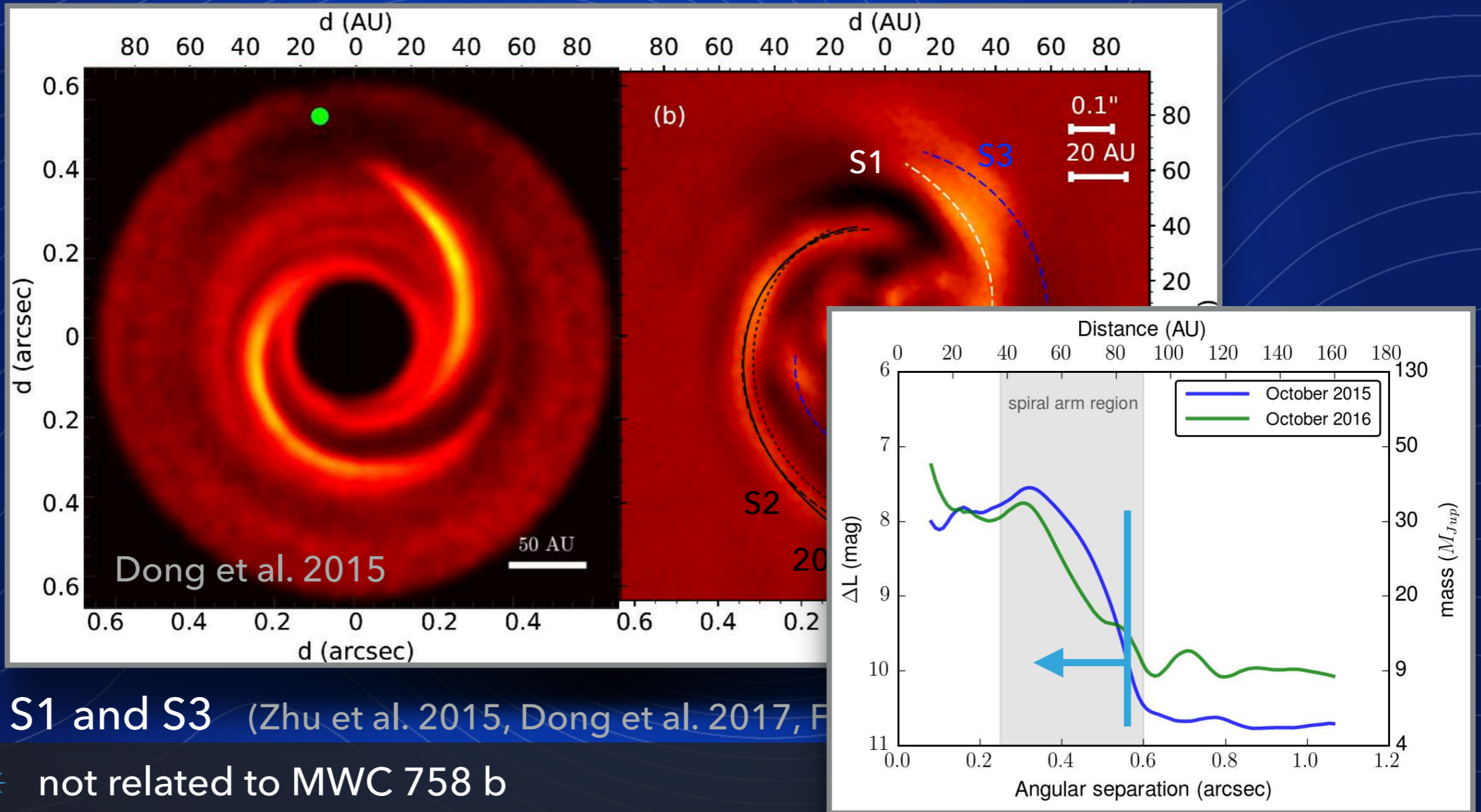
## SPIRAL ARMS: PURE GEOMETRICAL FITTING



S1 and S3 (Zhu et al. 2015, Dong et al. 2017, Fung et al. 2015):

- \* not related to MWC 758 b
- \* consistent with primary and secondary arm from external companion
- \* given the separation between them:  $M_p = 7 \pm 1 M_{Jup}$  located at  $< 100$  au

# SPIRAL ARMS: PURE GEOMETRICAL FITTING



S1 and S3 (Zhu et al. 2015, Dong et al. 2017, F)

- \* not related to MWC 758 b
- \* consistent with primary and secondary arm from external companion
- \* given the separation between them:  $M_p = 7 \pm 1 M_{Jup}$  located at  $< 100$  au

## SUMMARY / CONCLUSIONS

- ▶ new L' band images of MWC 758 reveal the presence of a third spiral arm and a bright point-like source south of the star
- ▶ the bright emission:
  - \* is most likely an embedded protoplanet at 20 AU
  - \* with mass  $m_p < 6 M_J$
  - \* accreting at a rate of  $10^{-7} - 10^{-9} M_\odot/\text{yr}$
- ▶ the spirals are most likely not related to MWC 758b
- ▶ S1 and S3
  - \* could be generated by a yet undetected companion of  $m_p = 7 \pm 1 M_J$  at  $< 80 \text{ au}$

FOLLOW-UP OBSERVATIONS WITH ALMA AND IN THE H ALPHA LINE, TOGETHER WITH NEW DEDICATED SIMULATIONS, WILL PROVIDE FURTHER INSIGHTS ON THE NATURE OF THE PROTOPLANET AND ITS CONNECTION WITH THE SPIRAL ARMS

thank you!